Impact of chiral hyperonic three-body forces on neutron stars

Isaac Vidaña (INFN, Catania)

on behalf of

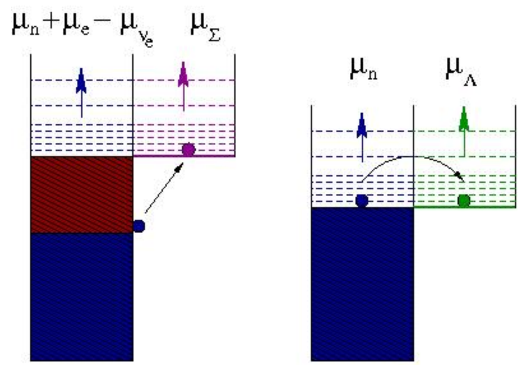
Domenico Logoteta (University of Pisa)

STRANEX: Recent progress and perspectives in STRANge EXotic atoms studies and related topics

October 21st - 25th 2019. ECT*, Trento (Italy)

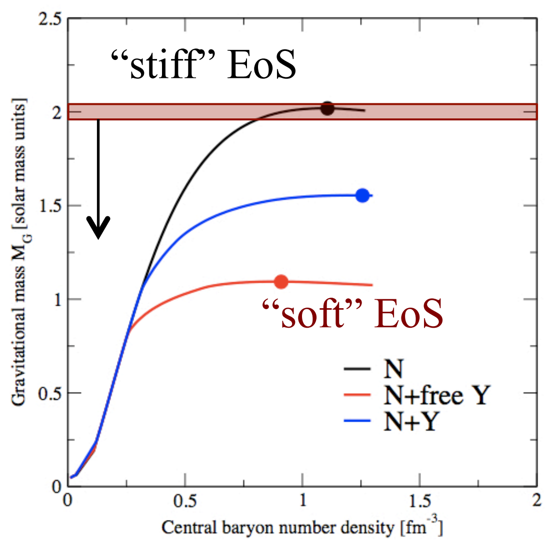
The Hyperon Puzzle: An Open Problem

Hyperons are expected to appear in the core of neutron stars at $\rho \sim (2\text{-}3)\rho_0$ when μ_N is large enough to make the conversion of N into Y energetically favorable.



But

The relieve of Fermi pressure due to its appearance $\longrightarrow$ EoS softer $\longrightarrow$ reduction of the mass to values incompatible with observation



Observation of $\sim 2\ M_\odot$ NS $\longrightarrow$ Any reliable EoS of dense matter should predict $M_{max}[EoS] > 2M_\odot$

Can hyperons be present in the interior of neutron stars in view of this new constraint ?

Possible Solutions to the Hyperon Puzzle

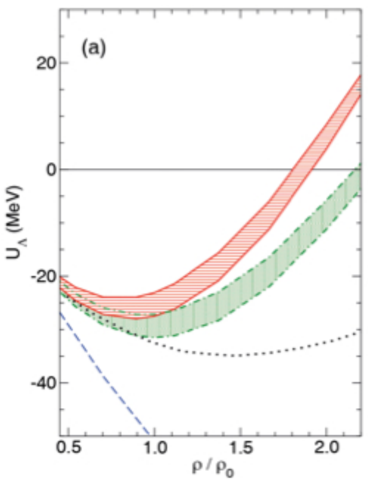
YN & YY

- YY vector meson repulsion

  ϕ meson coupled only to hyperons yielding strong repulsion at high ρ

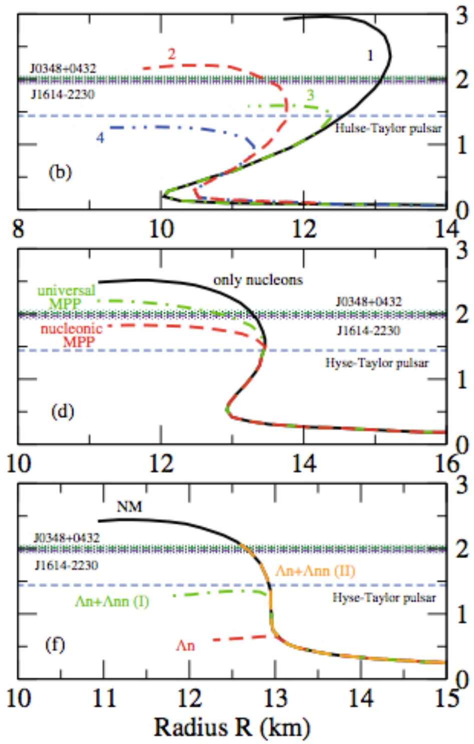
- Chiral forces

  YN from χEFT predicts Λ s.p. potential more repulsive than those from meson exchange



Hyperonic TBF

Natural solution based on the known importance of 3N forces in nuclear physics



Quark Matter

Phase transition to deconfined QM at densities lower that hyperon threshold

To yield $M_{max} > 2M_\odot$ QM should be

- significantly repulsive to guarantee a stiff EoS
- attractive enough to avoid reconfinement

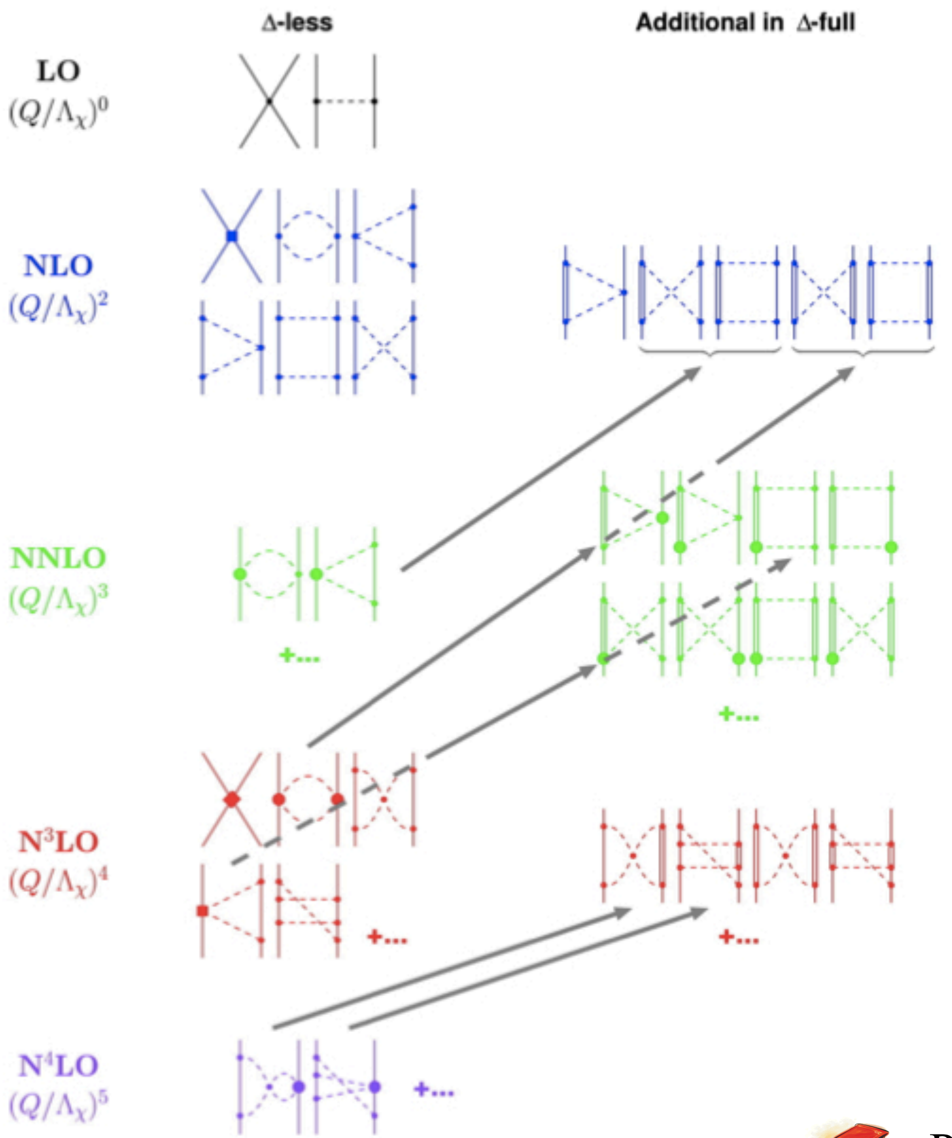
Short Summary of the Talk

- Study of the effects of NNΛ force on neutron stars
  - NN at N^3LO inχEFT including Δ isobar in inermediate states of NN scattering
  - NNN at N^2LO in χEFT
  - NΛ from meson-exchange (Nijmegen group). Weak point of the work
  - NNΛ derived by the Juelich-Munich-Bonn group in χEFT at N^2LO
  - EoS & NS structure derived within the BHF approach
- Inclusion of NNΛ force leads to an EoS stiff enough such that he resulting NS maximum mass is compatible with current observations but the model contains only N, leptons & Λ's
- We have NOT SOLVED the hyperon puzzle but have taken an additional step towards its solution

In collaboration with: Ignazio Bombaci (University of Pisa)

For details see:  arXiv:1906.11722. EpJA (in production)

NN interaction



- ✧ Local chiral potential derived in coordinate space
- ✧ Δ isobar included in intermediate NN scattering
- ✧ A gaussian regulator for short range contributions

$$\frac{1}{\pi^{3/2}R_s^3}\exp\left[-r^2/R_s^2\right], \quad R_S = 0.7\,fm$$

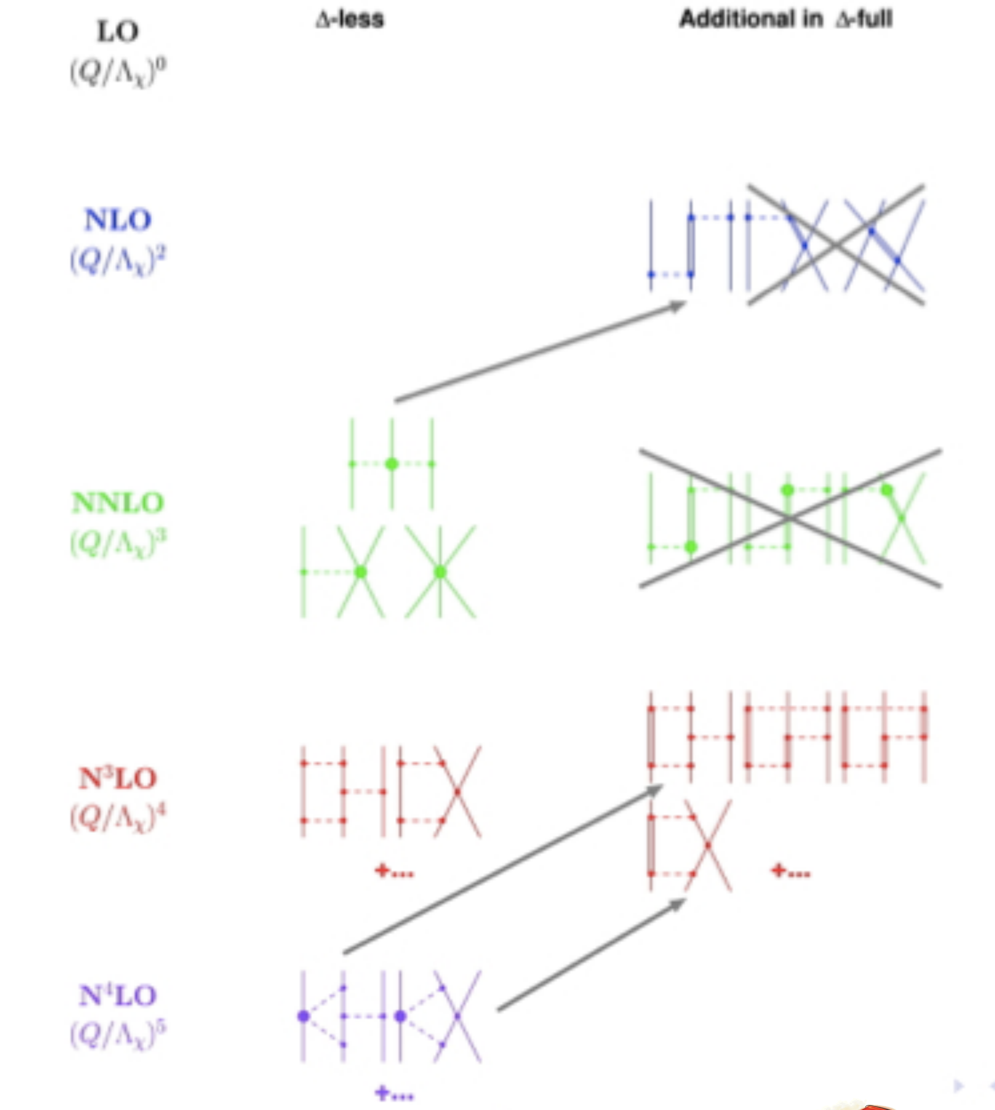
- ✧ Long range contributions regularized with

$$1-\left(\left(\frac{r}{R}\right)^6 \exp\left(\frac{2r}{R_l}-2\right)+1\right)^{-1}, \quad R_l = 1\,fm$$

- ✧ Operatorial structure similar to the Av18 (but not the same)

Piraulli et al., PRC 94, 054007 (2016)

NNN interaction



- First TNF contribution at NLO
- Δ isobar included
- Gaussian regulator used

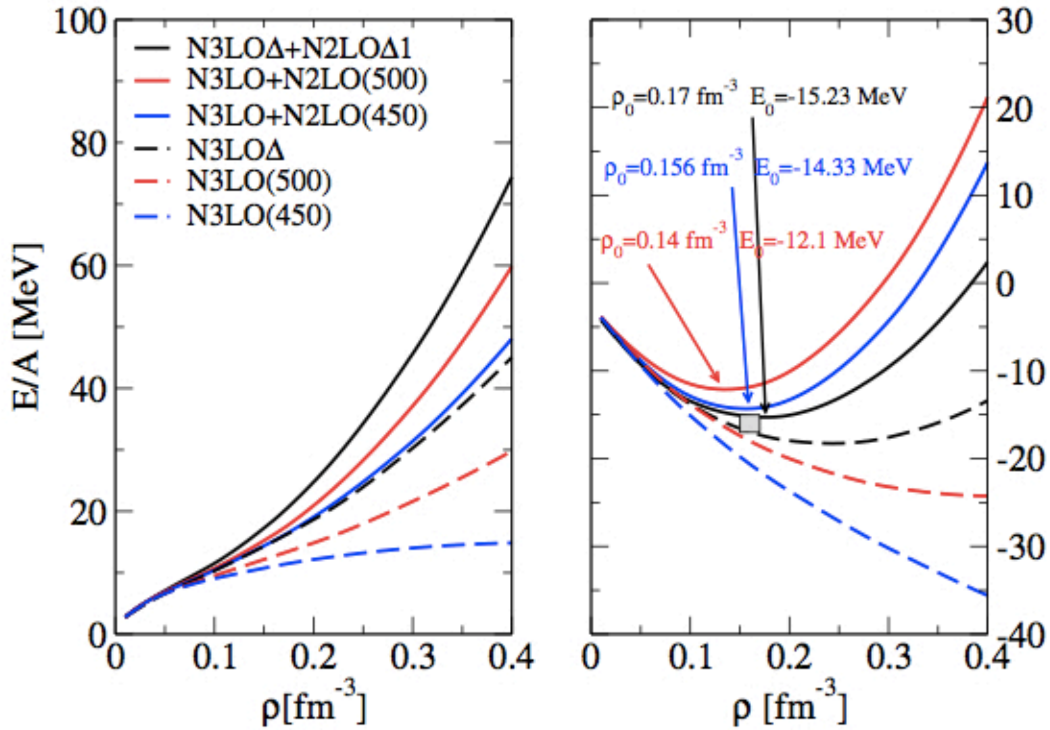
$$\frac{1}{\pi^{3/2}R_s^3}\exp\left[-r^2/R_s^2\right],\quad R_S = 0.7\,fm$$

- 5 LEC:
  - c_1, c_3, c_4 fixed from the two-body interaction
  - c_E & c_D fixed from observables of few body nuclear systems (^3H) or to reproduce the saturation point (this work)

Piraulli et al., PRC 94, 054007 (2016)

Nuclear Matter EoS

PNM SNM



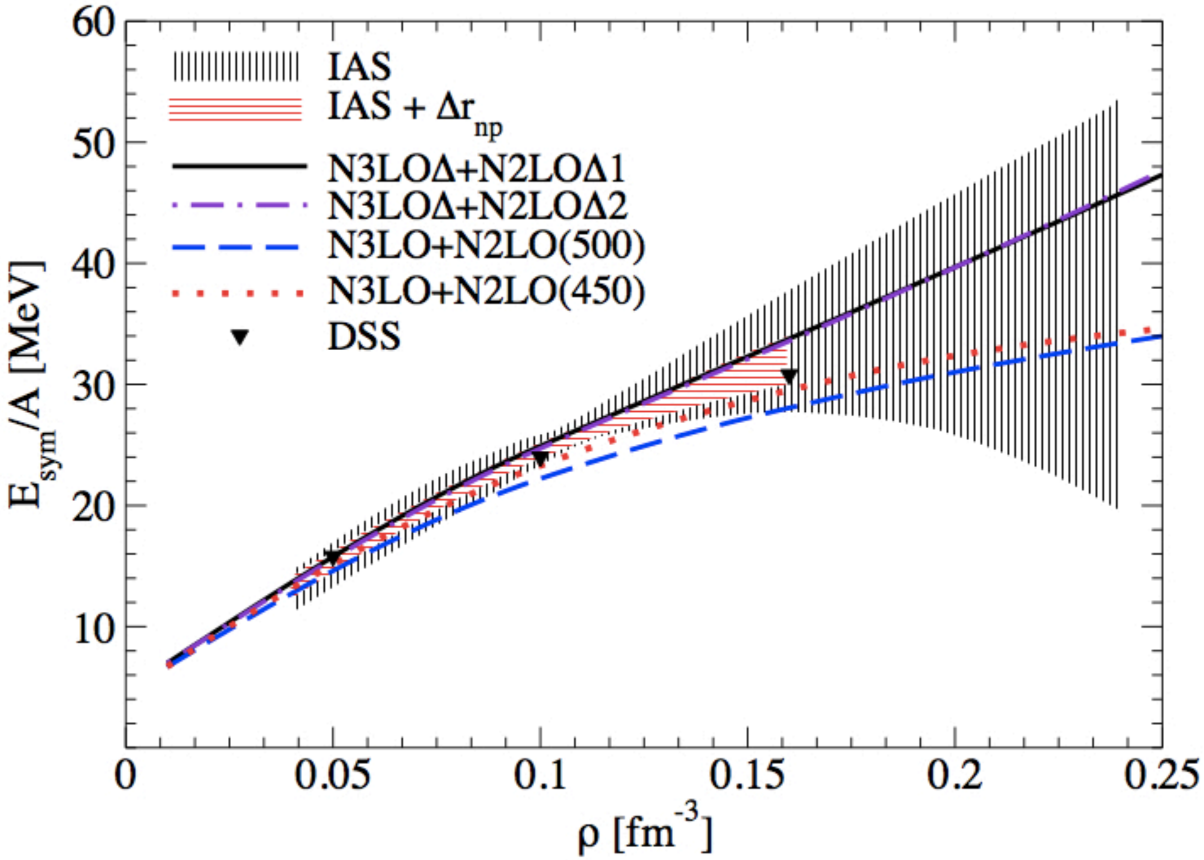
- ✧ Significant improvement of SNM saturation point due to TNF
- ✧ Explicit inclusion of the Δ isobar diminishes the strength of the TNF needed to obtain a good saturation point

Nuclear Matter Properties at Saturation

Model	ρ_0(fm^{-3})	E/A (MeV)	E_{sym} (MeV)	L (MeV)	K_∞ (MeV)
N3LOΔ+N2LOΔ1	0.171	-15.23	35.39	76.0	190
N3LOΔ+N2LOΔ2	0.176	-15.09	36.00	79.8	176
N3LO+N2LO(500)	0.135	-12.12	25.89	38.3	153
N3LO+N2LO(450)	0.156	-14.32	29.20	39.8	205

Logoteta et al., PRC 94, 064001 (2016)

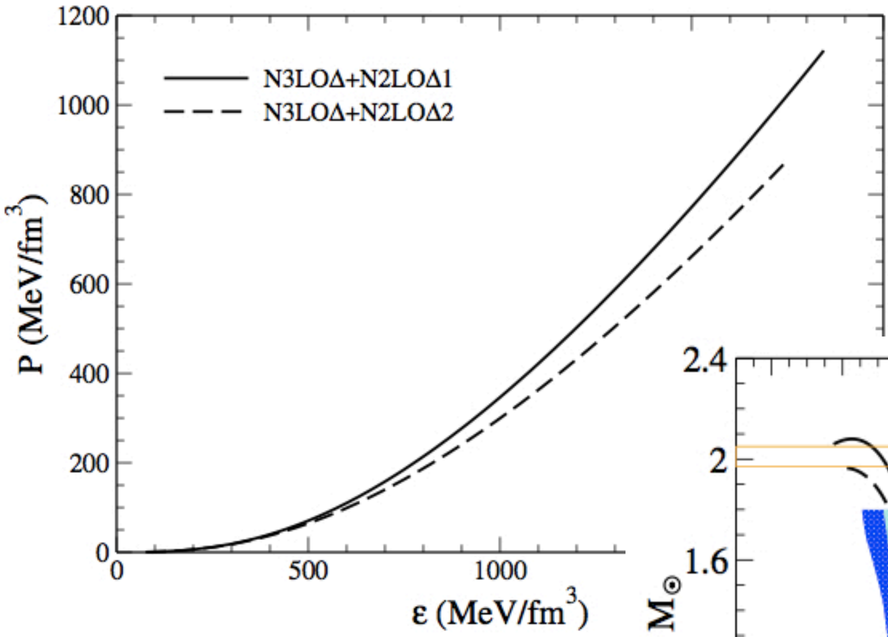
Nuclear Symmetry Energy

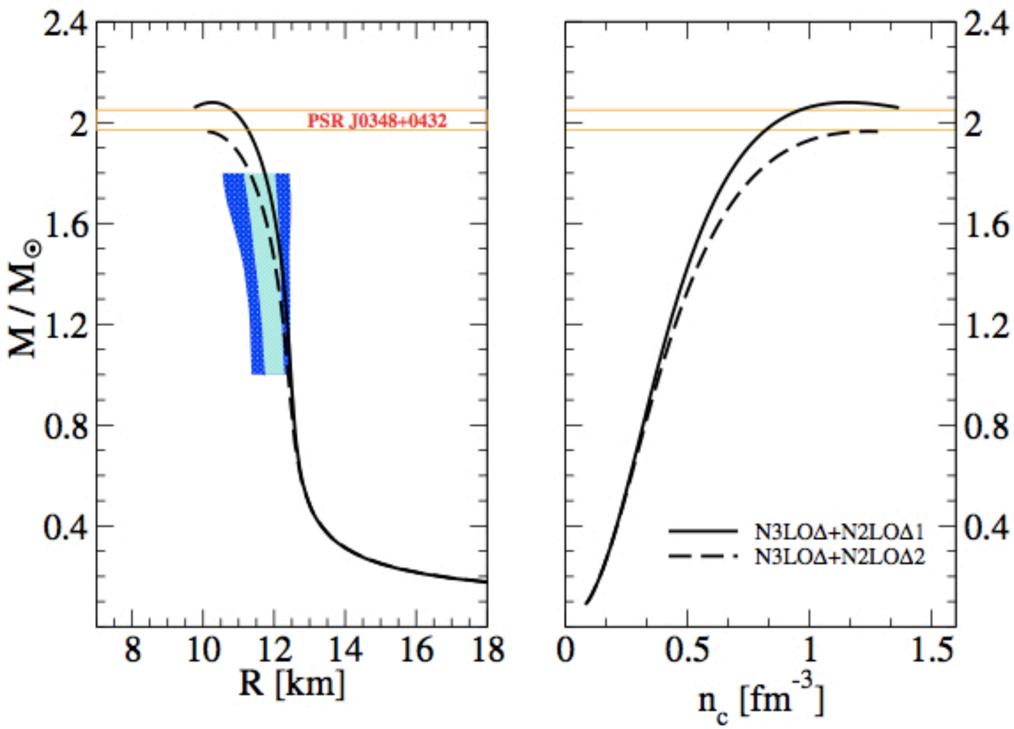


Good agreement with experimental constraints from Isobaric Analog States & neutron skin thinkness

Logoteta et al., PRC 94, 064001 (2016)

NS EoS & Structure at N^{3}LOΔ+N^{2}LOΔ





Good agreement with with recent observations of $2M_\odot$ NS

Logoteta et al., A&A 609, A128 (2018)

NΛ interaction

✧ NY & YY interactions in χEFT derived by the Juelich-Bonn-Munich group

- NY at

  LO: Polinder et al., NPA 779, 244 (2006)

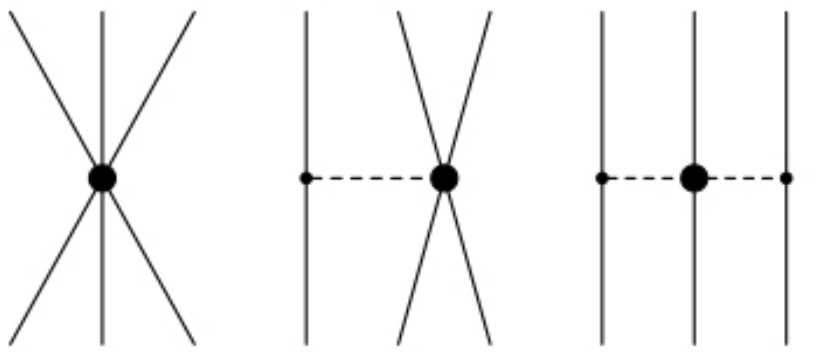
  NLO: Haidenbauer et al., NPA 915, 24 (2013)

  Haidenbauer et al., arXiv:1906.11681

- YY at NLO: Haidenbauer et al., NPA 954, 273 (2016)

WEAK POINT OF THE PRESENT WORK: unfortunately at present we do not have at our disposal this interaction and, therefore, instead we use a NY meson-exchange interaction from the Nijmegen group (NSC97a & NSC97f)

NNΛ interaction



(π exchange expected to be dominant, heavy meson (K,η) exchange absorbed into contact terms)

- ✧ First contributions to NNY appear in χEFT at N^2LO
- ✧ Leading terms at N^2LO
  - ✓ Three-baryon contact terms
  - ✓ One & two-meson exchange

- ✧ LEC estimated through decouplet saturation. Only one LEC H' remains a free parameter. A value of $H'=\pm 1/f_\pi$ where $f_\pi=93$ MeV has been considered by Petschauer et al.
- ✧ In this work $H'=\beta/f^2_\pi$ with β a parameter fixed to reproduce $U_\Lambda(k=0) \sim -28 - 30$ MeV in SNM at saturation ⟶ Models $NN\Lambda_1$ & $NN\Lambda_2$
- ✧ A non-local regulator $\exp(-(p^4+p'^4)/\Lambda^4)$ with Λ=500 MeV is used

Petschauer et al., PRC93, 014001 (2016)

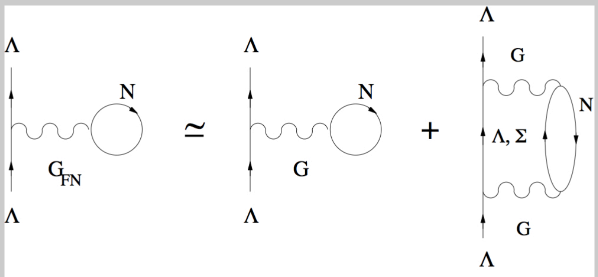
Effect of NNΛ interaction on hypernuclei

Before analyzing the effect of NNΛ interaction on NS we consider its role on hypernuclear strucure. To such end

A perturbative many-body approach is employed to obtain the Λ self-energy in finite nuclei from which the Λ s.p. bound states can be obtained

$$G_{NM} = V + V\left(\frac{Q}{E}\right)_{NM} G_{NM} \longrightarrow G_{FN} = G_{NM} + G_{NM}\left[\left(\frac{Q}{E}\right)_{FN} - \left(\frac{Q}{E}\right)_{NM}\right] G_{FN}$$

Λ irreducible self-energy in finite nuclei



(for details of the method see Vidaña's talk on charmed nuclei on Wednesday)

Effect of NNΛ interaction on hypernuclei

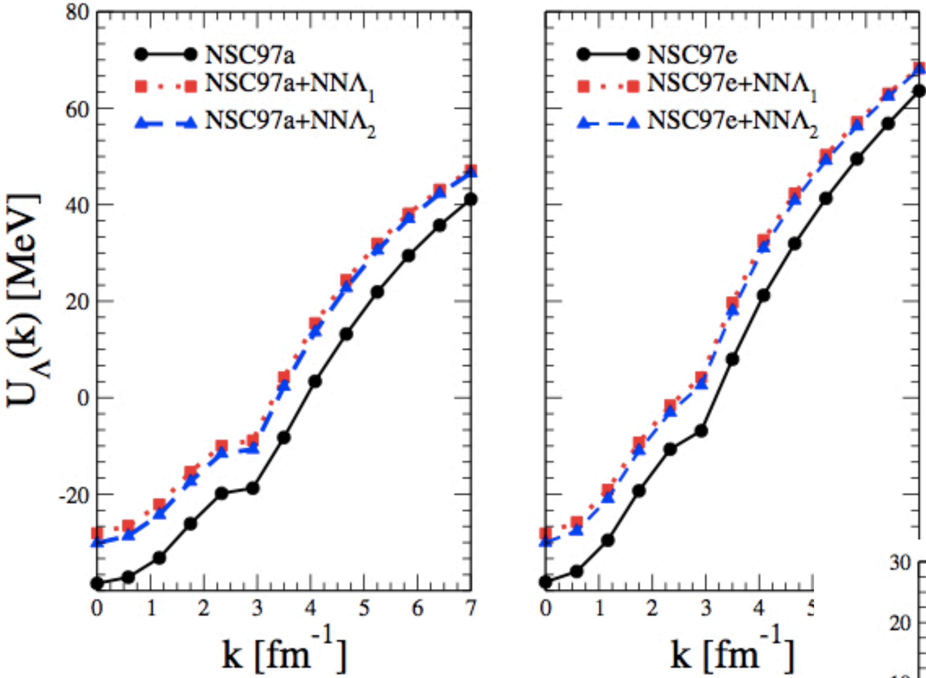
Λ separation energy in ${}^{41}_{\Lambda}$Ca, ${}^{91}_{\Lambda}$Zr & ${}^{209}_{\Lambda}$Pb

	${}^{41}_{\Lambda}$Ca	${}^{91}_{\Lambda}$Zr	${}^{209}_{\Lambda}$Pb
NSC97a	23.0	31.3	38.8
NSC97a+NNΛ_1	14.9	21.1	26.8
NSC97a+NNΛ_2	13.3	19.3	24.7
NSC97e	24.2	32.3	39.5
NSC97e+NNΛ_1	16.1	22.3	27.9
NSC97e+NNΛ_2	14.7	20.7	26.1
Exp.	18.7(1.1)*	23.6(5)	26.9(8)

- ✧ We consider only hypernuclei described as a closed shell nuclear core + a Λ sitting in a s.p. state. Comparison with the closest hypernucleus for which exp. data is available
- ✧ Inclusion of NNΛ improves the agreement with data for ${}^{91}_{\Lambda}$Zr & ${}^{209}_{\Lambda}$Pb.
- ✧ NNΛ predict too much repulsion in the case ${}^{41}_{\Lambda}$Ca & lighter hypernuclei
- ✧ No refit of any parameter has been done to perform these calculatios

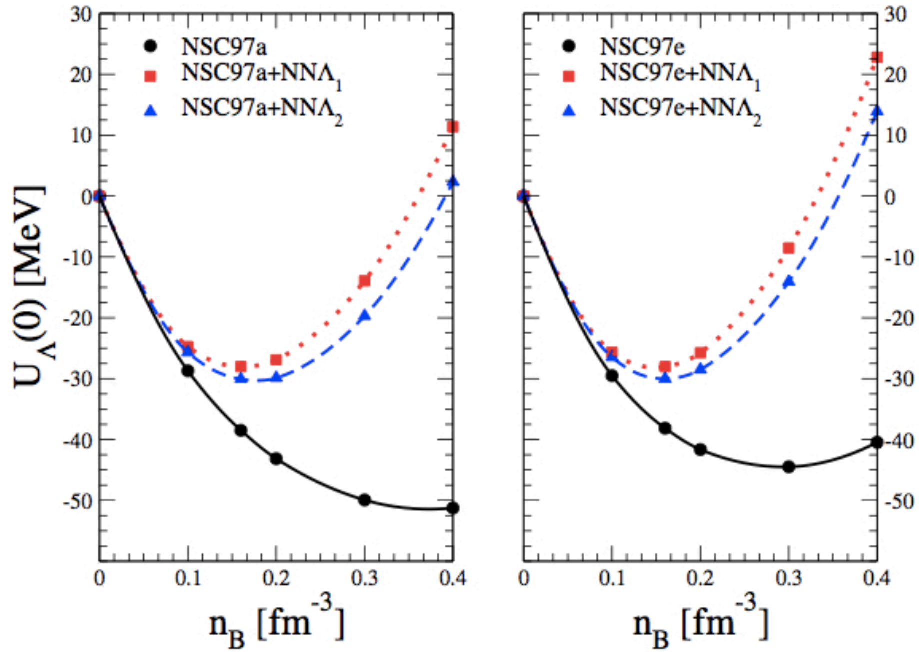
* Taken from Pile et al., PRL 66, 2585 (1991). Not included in the recent Gal et al., RMP 88, 035004 (2016)

Λ single-particle potential in SNM at ρ_0

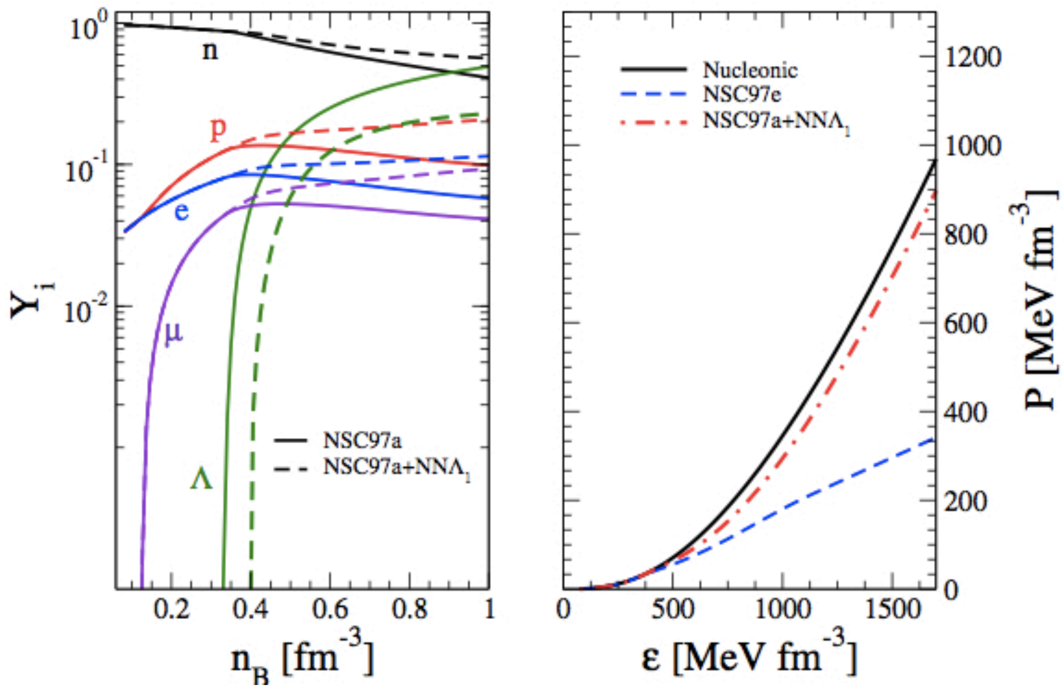


- ✧ NSC97a & NSC97e (with no NNΛ) predict $U_\Lambda(0) \sim -40$ MeV.
- ✧ NSC97a predict more attraction than NSC97e even when NNΛ is included
- ✧ NNΛ induces as expected repulsion

- ✧ $U_\Lambda(0)$ very deep when only two-body forces are considered
- ✧ NNΛ induces repulsion for densities larger than ~0.1 fm^{-3} & shift the minimum to ~0.16 fm^{-3}
- ✧ NNΛ negligible in the low density region



Neutron star matter EoS & Composition



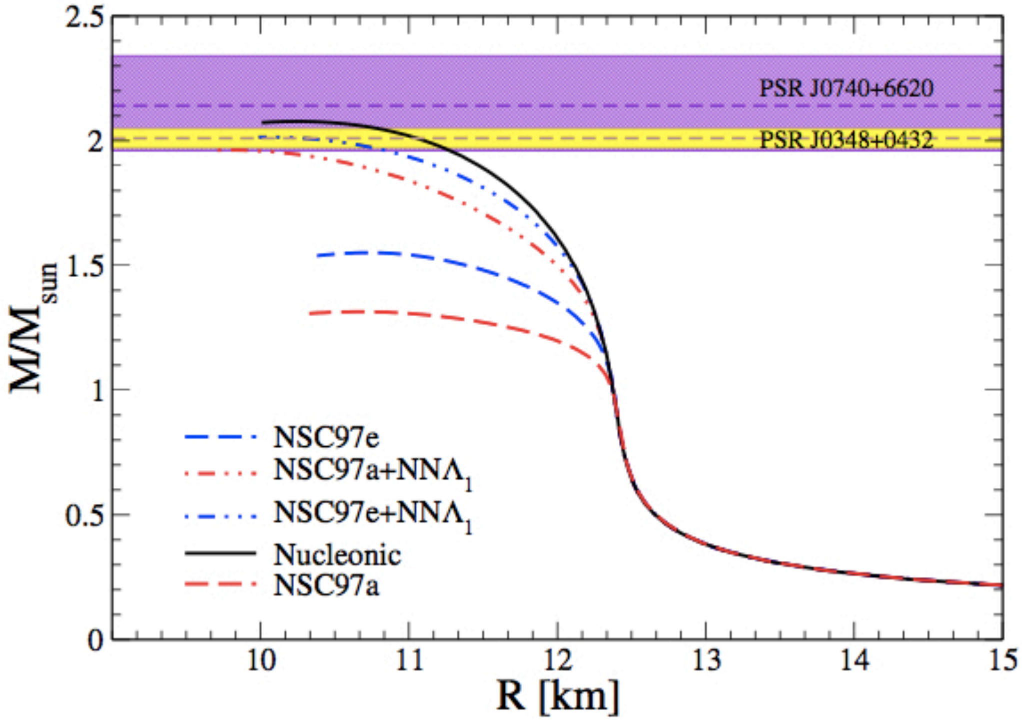
Only n, p, e⁻, μ⁻ & Λ included

- ✓ First exploratory work. We are interested on the role of NNΛ
- ✓ More complete study requires the inclusion of other hyperons (work on progress)

✧ The effect of NNΛ interaction is twofold:

- ✓ Shift the onset of the Λ to slighly larger baryon density
- ✓ Strong reduction of the amount of Λ's at large baryon densities with the consequent stiffening of the EoS compared to the case in which the NNΛ is not included ⟶ important consequences for NS mass (M_{max} increases)

Neuron star Properties



	$M_{max}(M_\odot)$	R (km)	n_c (fm^{-3})
Nucleonic	2.08	10.26	1.15
NSC97a	1.31	10.60	1.40
NSC97a+NNΛ_1	1.96	9.80	1.30
NSC97a+NNΛ_2	1.97	9.87	1.28
NSC97e	1.54	10.81	1.18
NSC97e+NNΛ_1	2.01	10.10	1.20
NSC97e+NNΛ_2	2.02	10.15	1.19

NS M_{max} compatible with the largest NS observed

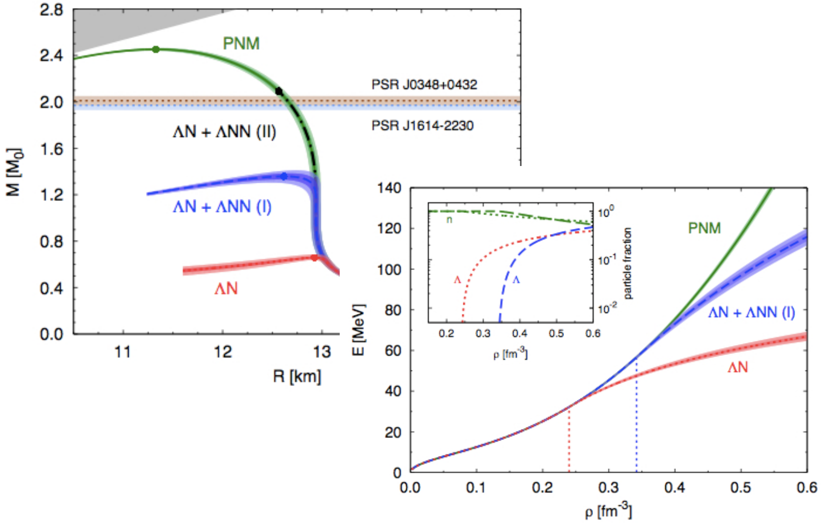
But

- ✓ We have ignored the presence of other hyperons in the NS interior that could change this conclusion
- ✓ Hypothetical repulsive NNY, NYY & YYY forces could lead to a similar conclusion

In view of this we CANNOT say that we have solved the hyperon puzzle

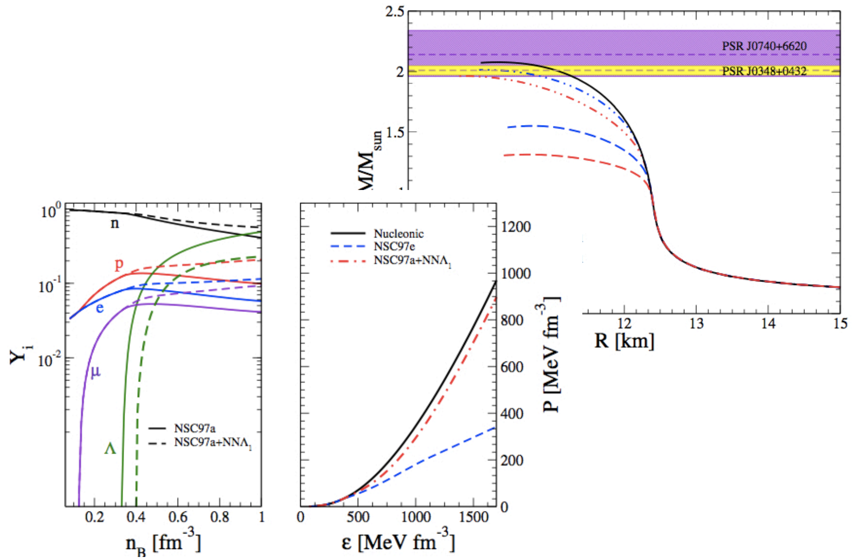
A comparison with QMC calculations

Lonardoni et al., PRL 114, 092301 (2016)



- ✧ NS matter described as mixture of neutrons & Λ's in β-equilibrium
- ✧ Simple interaction models: Av8'+UIX (nn,nnn) & Bodmer-Usmani (NΛ) potential, NNΛ (2π exchange + phenomenological repulsive term)
- ✧ The only NNΛ able to give $2M_\odot$ lead to the total disappearance of Λ in NS , but this in in fact just pure neutron matter

This work



- ✧ NS matter described as a mixture of n, p, e$^-$, μ$^-$ & Λ's in β-equilibrium
- ✧ χEFT (NN, NNN, NNΛ) + meson-exchange (NY)
- ✧ Even if the concentration of Λ's is strongly reduced they are still present in the interior of a $2M_\odot$ NS

Summary & Conclusions (Again)

✧ Study of the effects of NNΛ force on neutron stars

- NN at N^3LO inχEFT including Δ isobar in inermediate states of NN scattering
- NNN at N^2LO in χEFT
- NΛ from meson-exchange (Nijmegen group). Weak point of the work
- NNΛ derived by the Juelich-Munich-Bonn group in χEFT at N^2LO
- EoS & NS structure derived within the BHF approach

✧ Inclusion of NNΛ force leads to an EoS stiff enough such that he resulting NS maximum mass is compatible with current observations but the model contains only N, leptons & Λ's

✧ We have NOT SOLVED the hyperon puzzle but have taken an additional step towards its solution

- ✧ You for your time & attention
- ✧ Domenico who could not come & forced me to work double this week