

Neural Networks for Experimental Nuclear Astrophysics

The $^{17}\text{O}(p,\gamma)^{18}\text{F}$ Resonance at 65 keV as a Test Case

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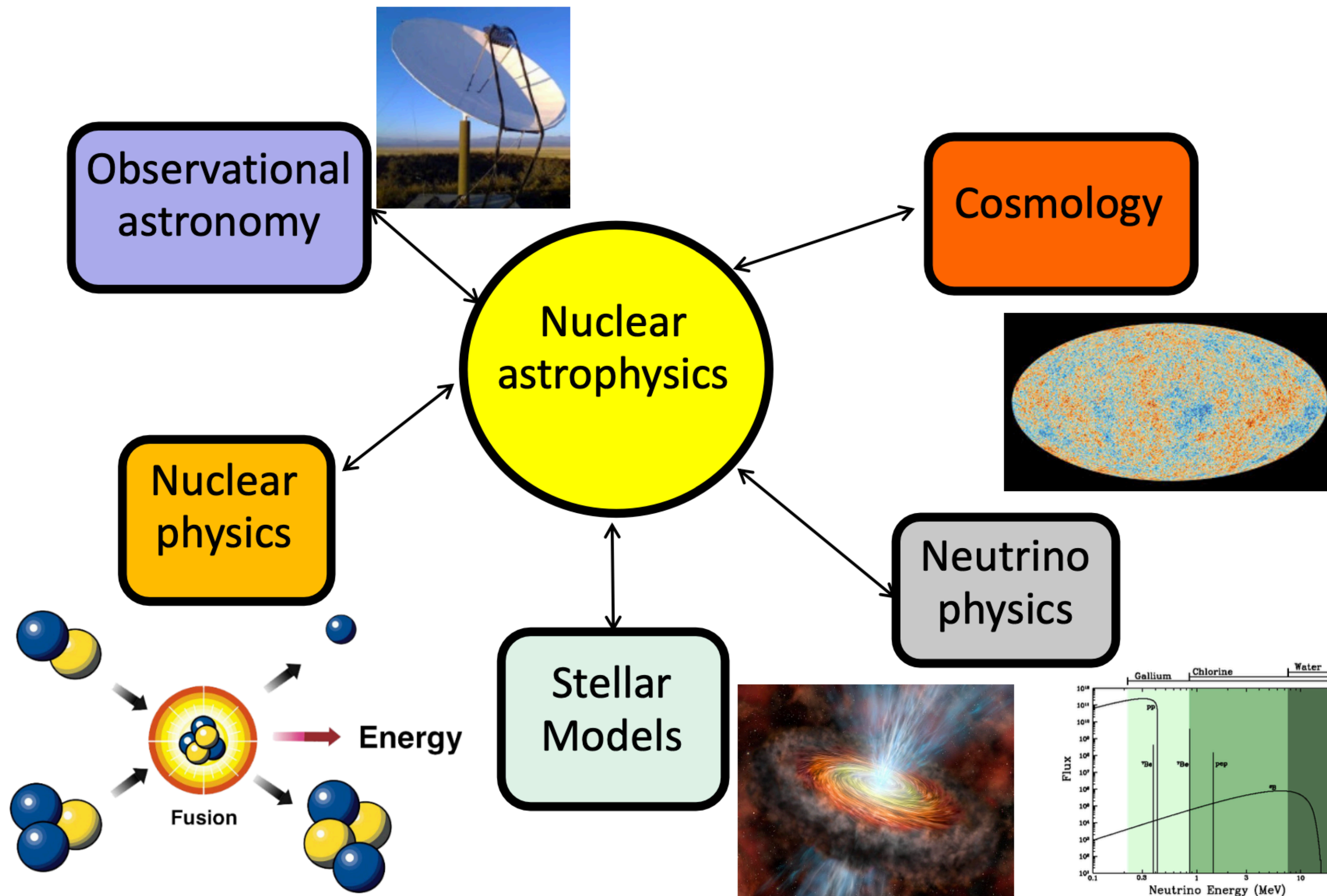
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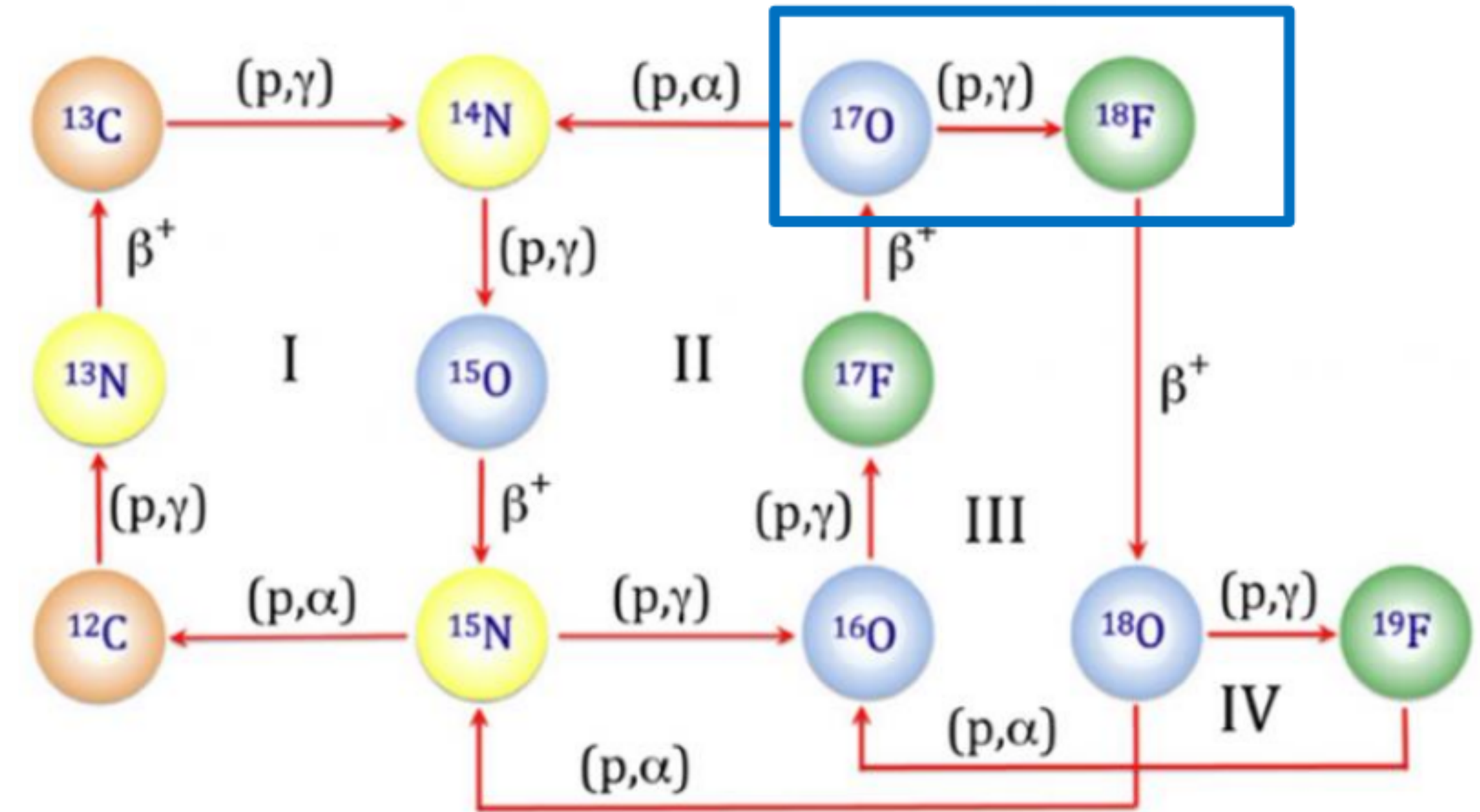


Nuclear Astrophysics



Astrophysical Motivation

- The $^{17}\text{O}(p,\gamma)^{18}\text{F}$ reaction takes part in the **CNO cycle**
- The CNO cycle is the predominant **energy production mechanism** for heavy stars ($M > 1.3 M_{\odot}$)
- The reaction rates of the CNO cycles determines the **abundances of elements** inside the **stellar core**



CNO Cycle

**Measuring Reaction
Cross Sections**

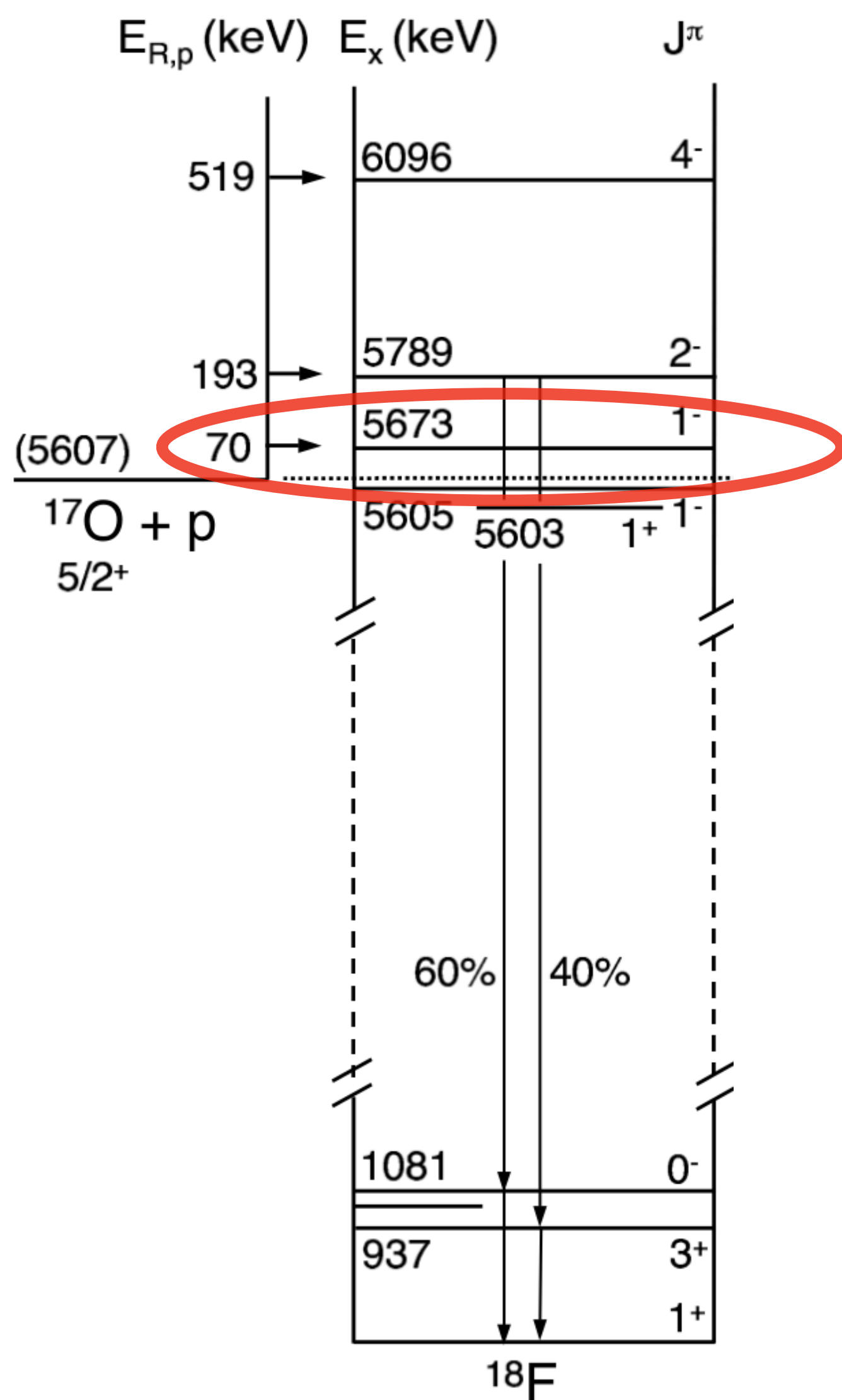


**Describing Stellar Evolution
and Nucleosynthesis**

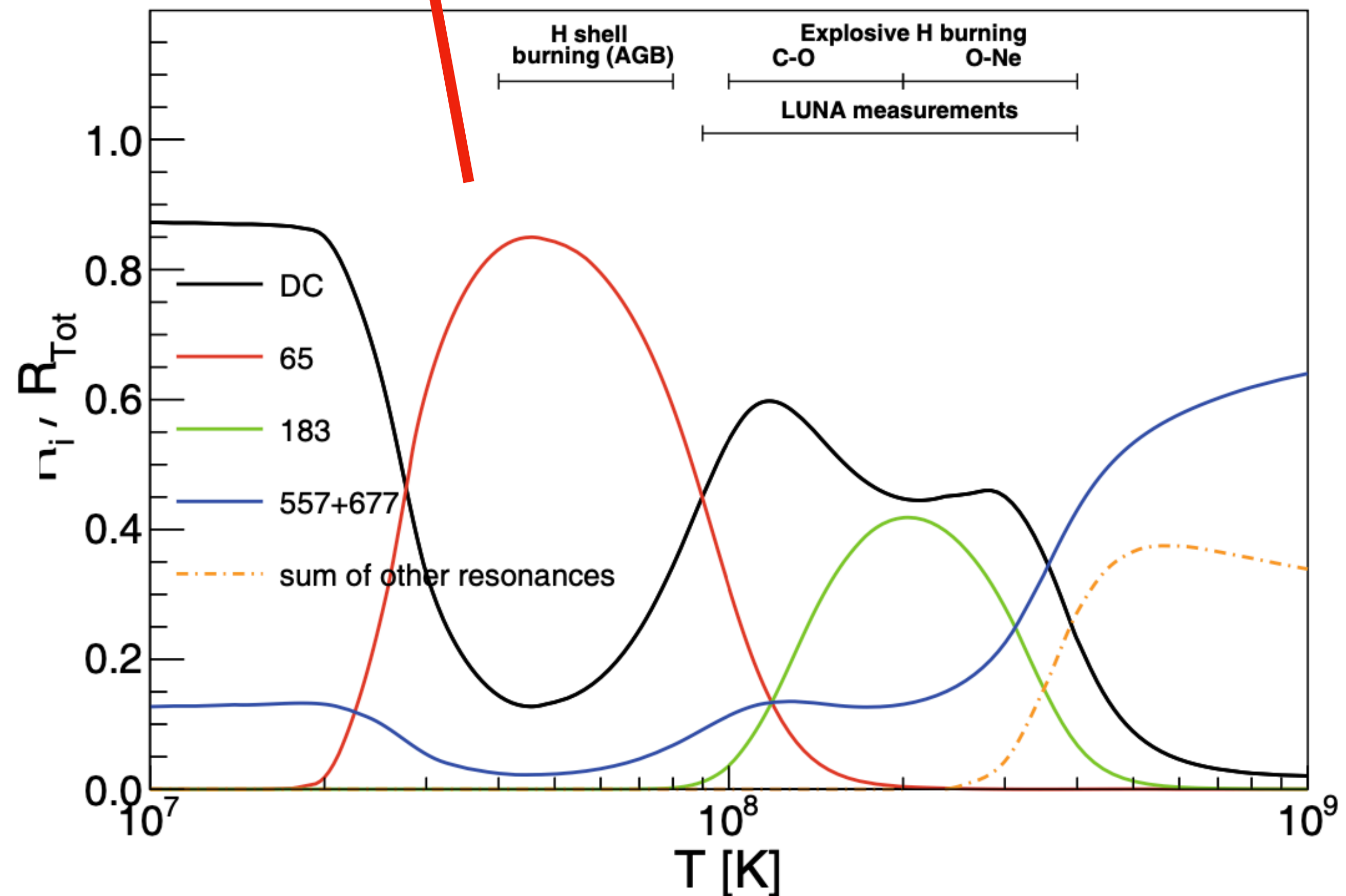


$^{17}\text{O}(p,\gamma)^{18}\text{F}$ Reaction

The reaction rate in the range of interest is dominated by the resonance at 65 keV!



Reaction Rate



65 keV Resonance

No direct measurements available

Estimated strength of the resonance: $\omega\gamma = (1.6 \pm 0.3) \times 10^{-11} \text{ eV}$ ($\omega\gamma \sim$ Cross Section)

—————→ Only **1 count per day** (considering a proton current of 100 μA)

- High detection efficiency ($\sim 50\%$)

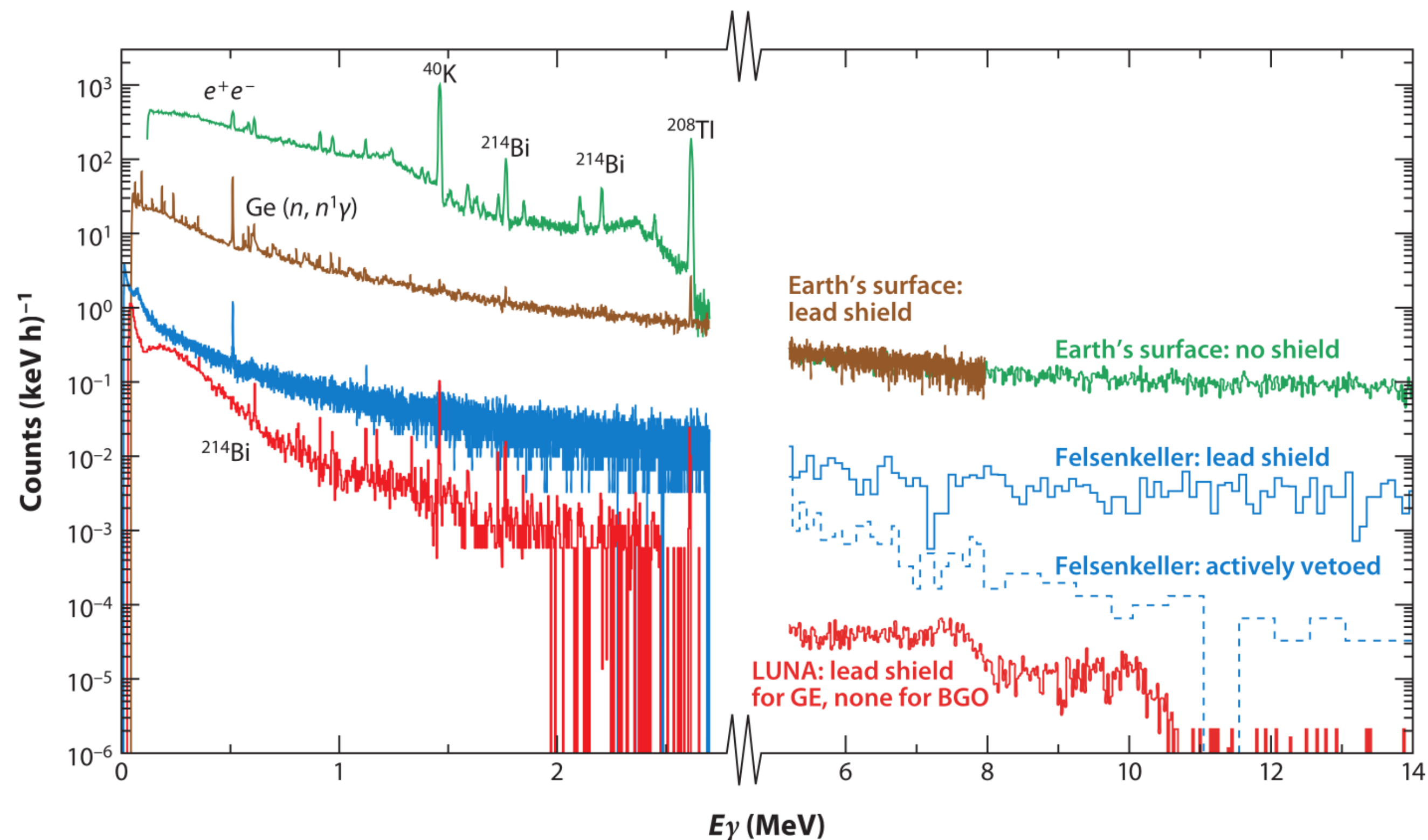
What we need:

- Environment with exceptional background reduction (< 1 event / day)
- Optimal signal / noise ratio



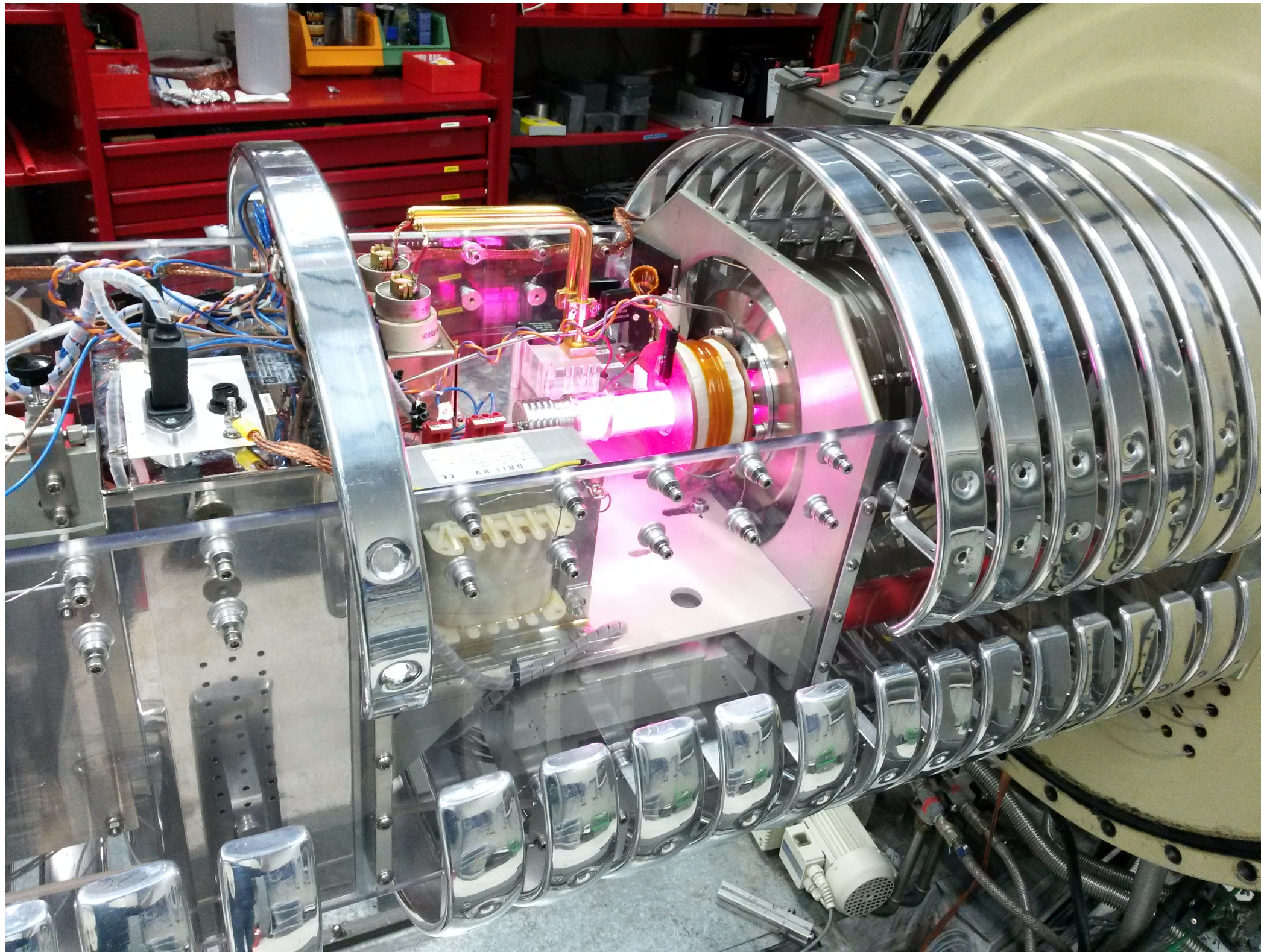
Laboratory for Underground Nuclear Astrophysics

- Located at **LNGS** facility under the **Gran Sasso mountain** in Abruzzo, Italy
- The cosmic ray flux reduced by **six orders of magnitude**



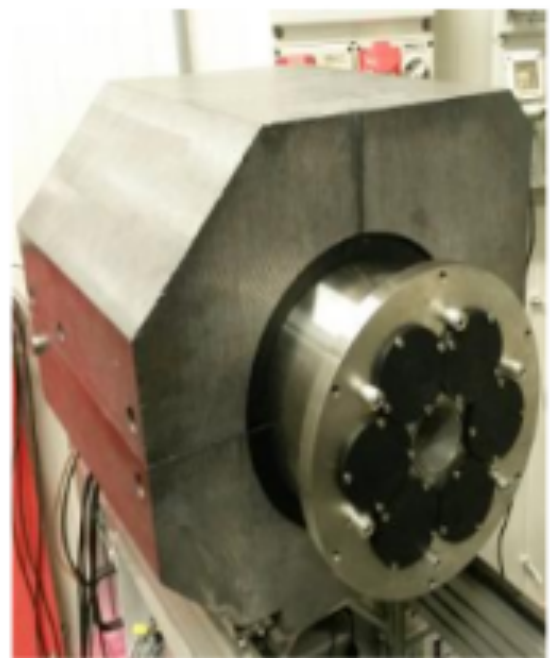
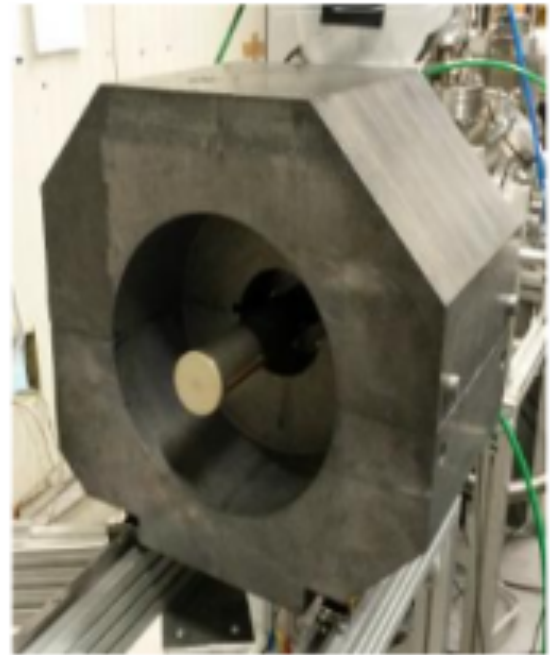
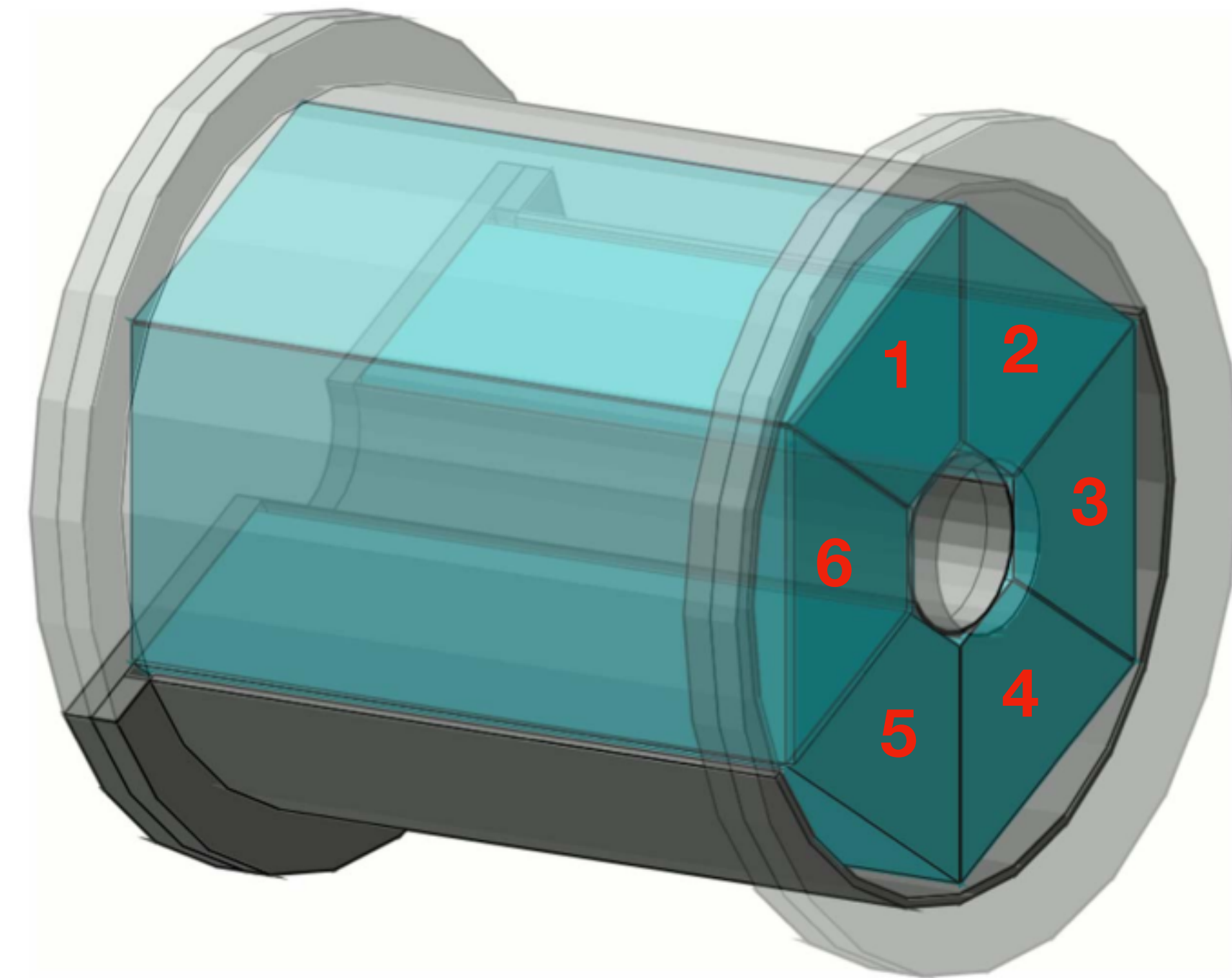
Experimental Setup - BGO Detector

LUNA 400kV Accelerator



- p beam up to $400\ \mu\text{A}$
- $E_p = 50 - 400\ \text{keV}$

BGO Detector

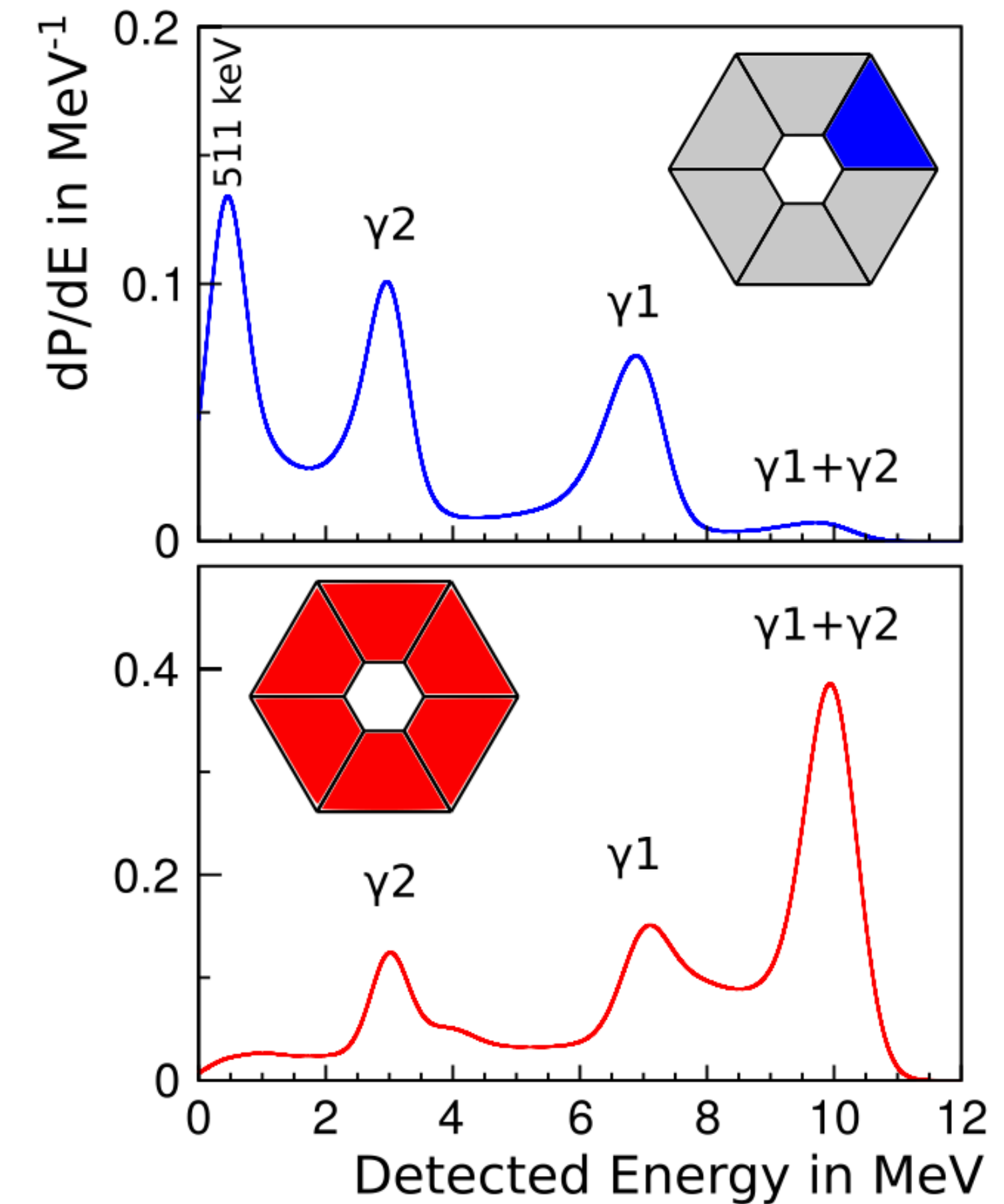
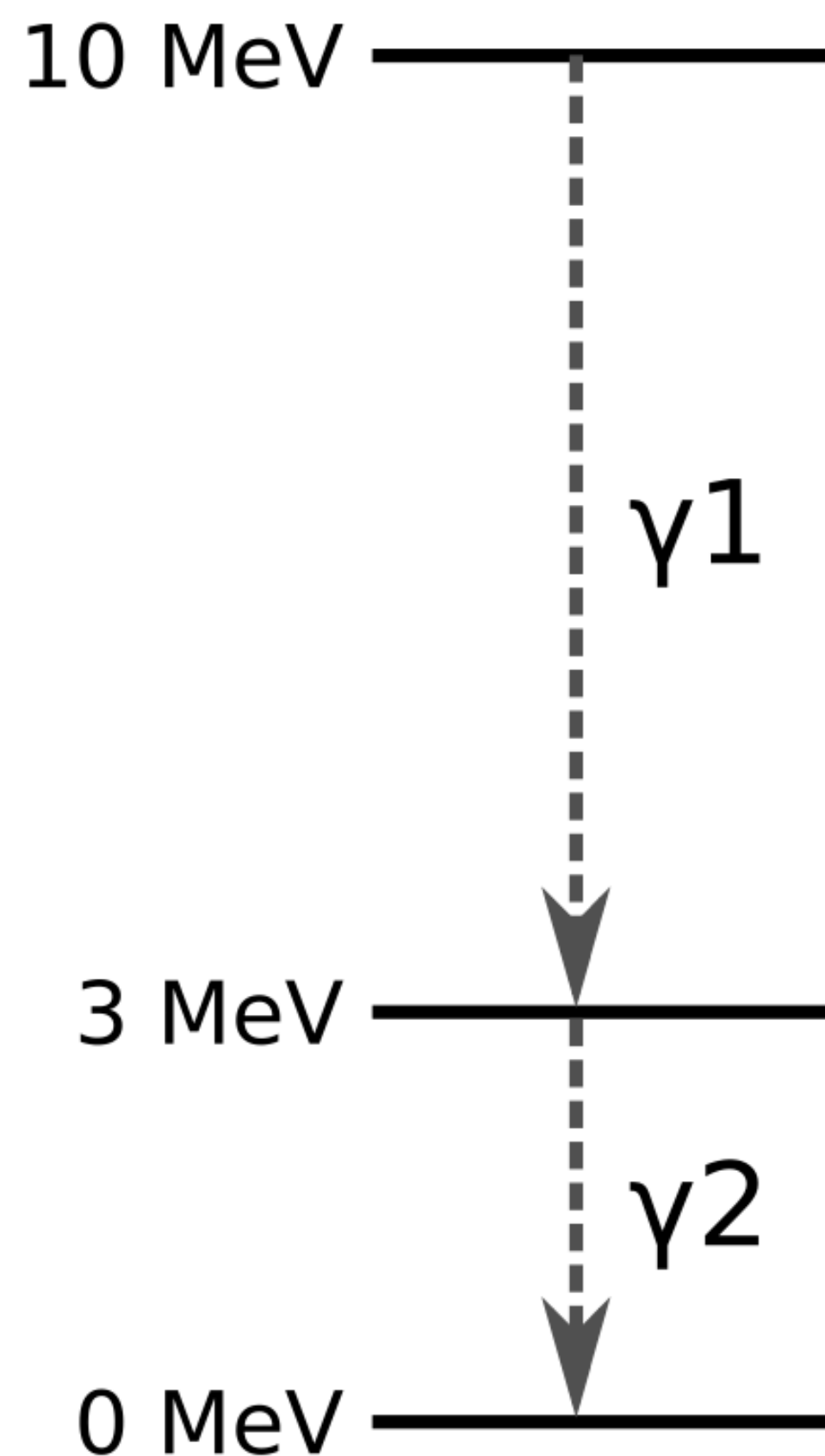


- Almost 4π geometry
- Segmented in 6 different crystals

Total Absorption Spectroscopy

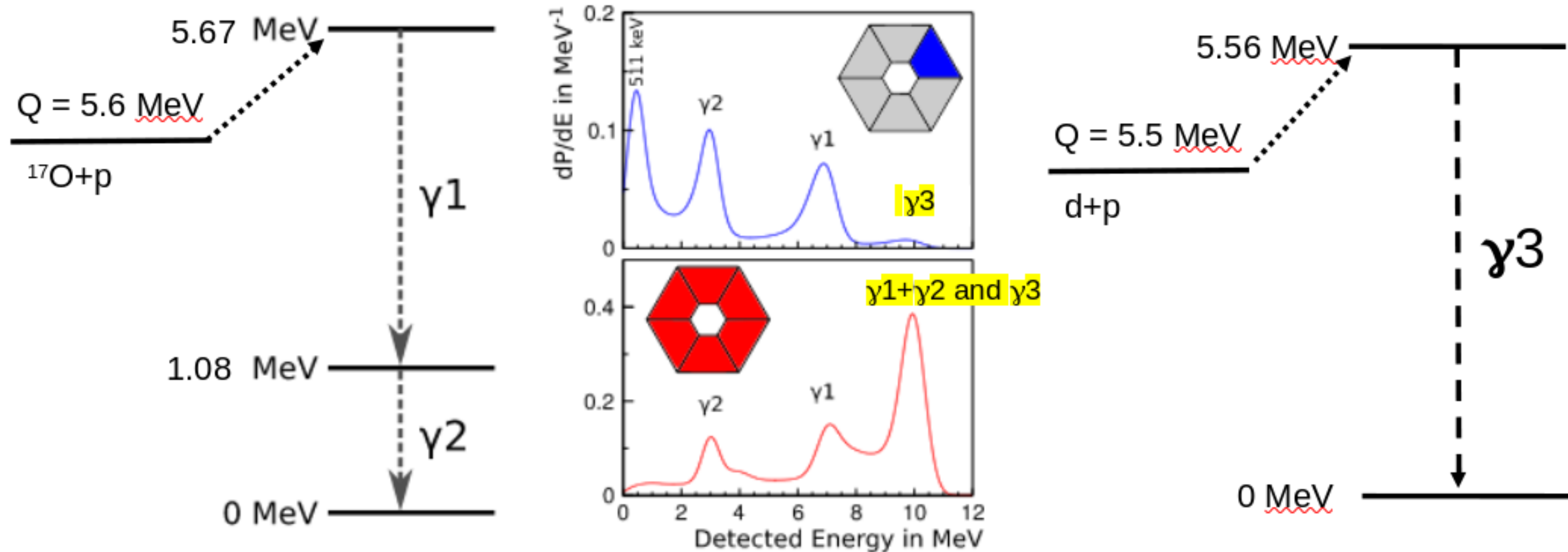
Idea

1. Detect **all the γ -rays** in coincidence
2. Construct the **sum γ -peak** by summing all the crystals
3. **Count the events** inside the sum γ -peak
4. Calculate the **cross section**



Deuteron Background

Problem: Ta backing (where ^{17}O is evaporated) contains deuterons...



Classic: gating the spectrum on the two γ -rays (thus reducing statistics)

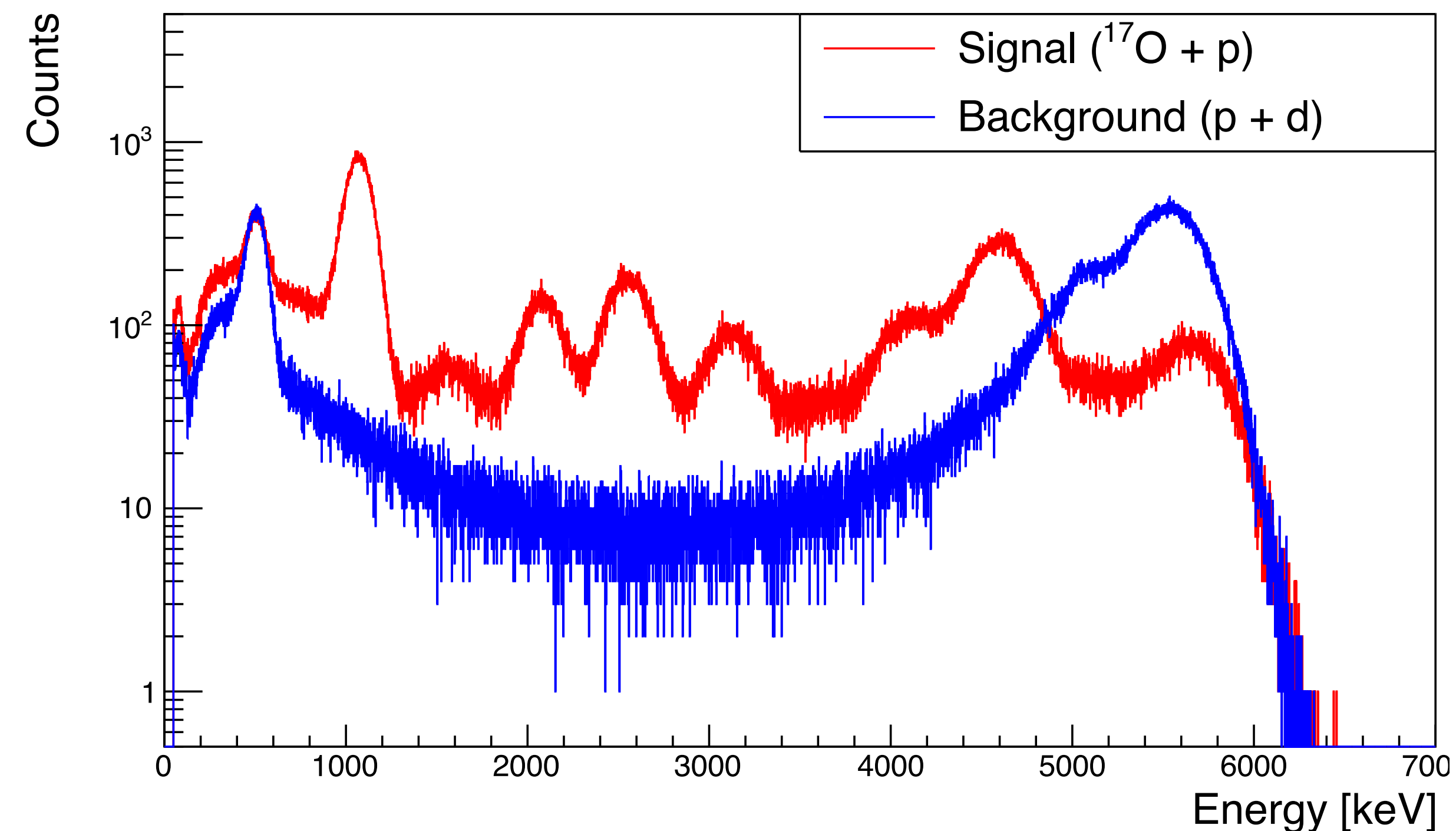
New: use Neural Networks to distinguish between $^{17}\text{O} + p$ and $p + d$ events

Training and Observed Data

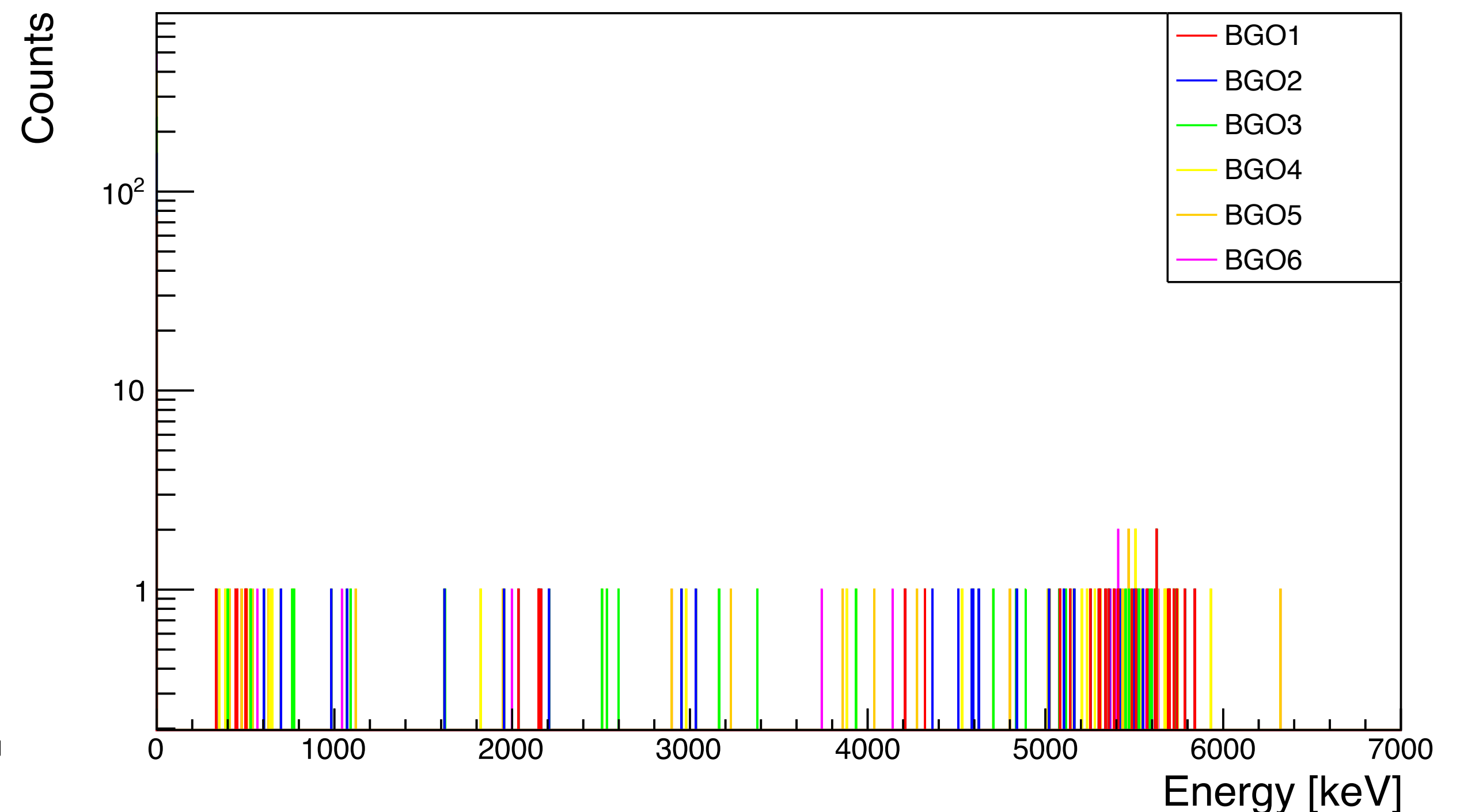
Geant4 simulations used as training data (**event by event**):

- Validated with calibration sources (**J. Phys. G: Nucl. Part. Phys. 50 045201**)
- Added random gain shifts to reproduce the BGO drifts

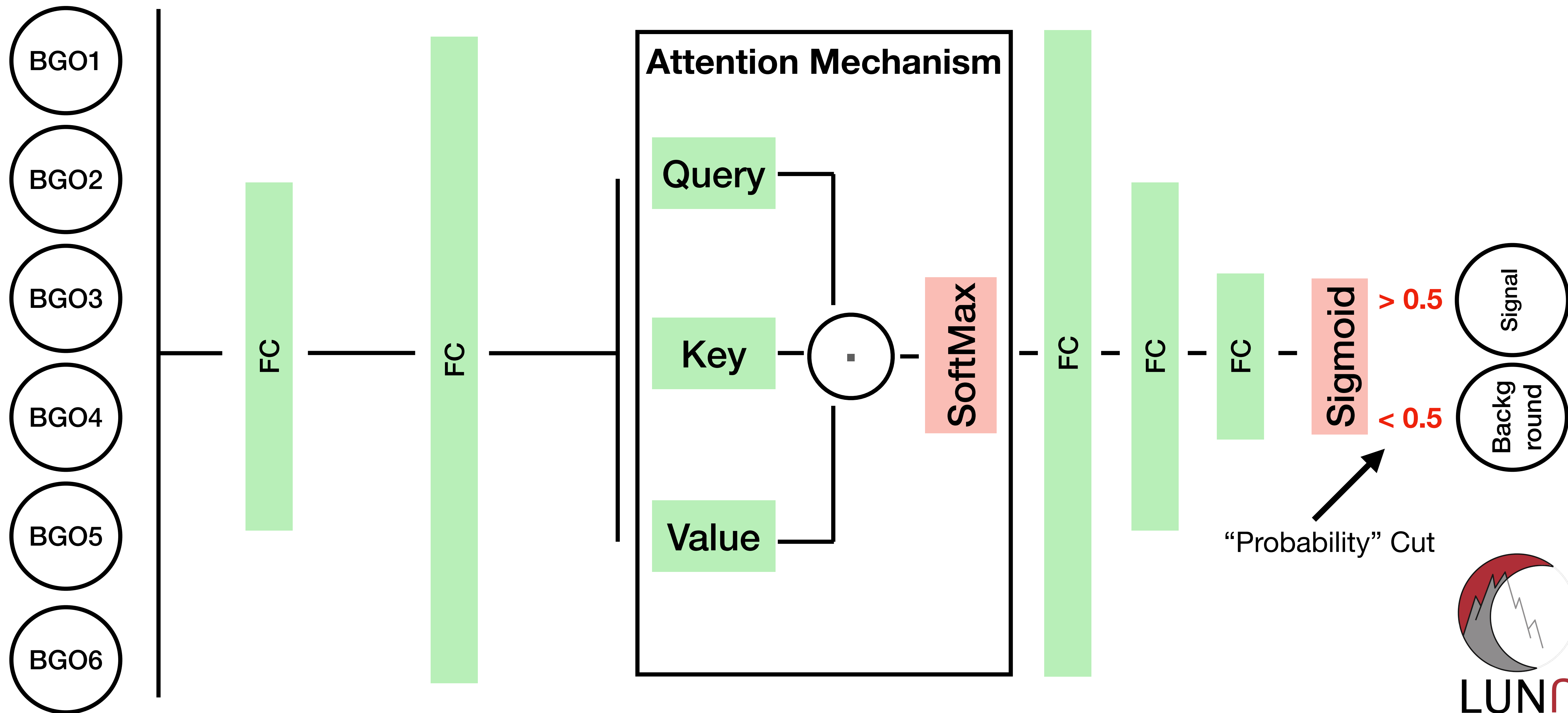
Training Data



Observed Data



Simple Neural Network Architecture



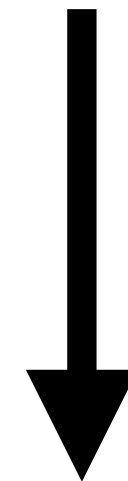
Classification

The accuracy of the classifier is defined as:

$$\text{Accuracy} = P(\text{ True Positives }) + P(\text{ True Negatives })$$

In order to extract the cross section it is important to estimate the **selection efficiency**, i.e. how many events we classify as signals respect the total signals:


$$\text{Selection Efficiency} = [1 - P(\text{ False Positives })] \times P(\text{ True Positives })$$

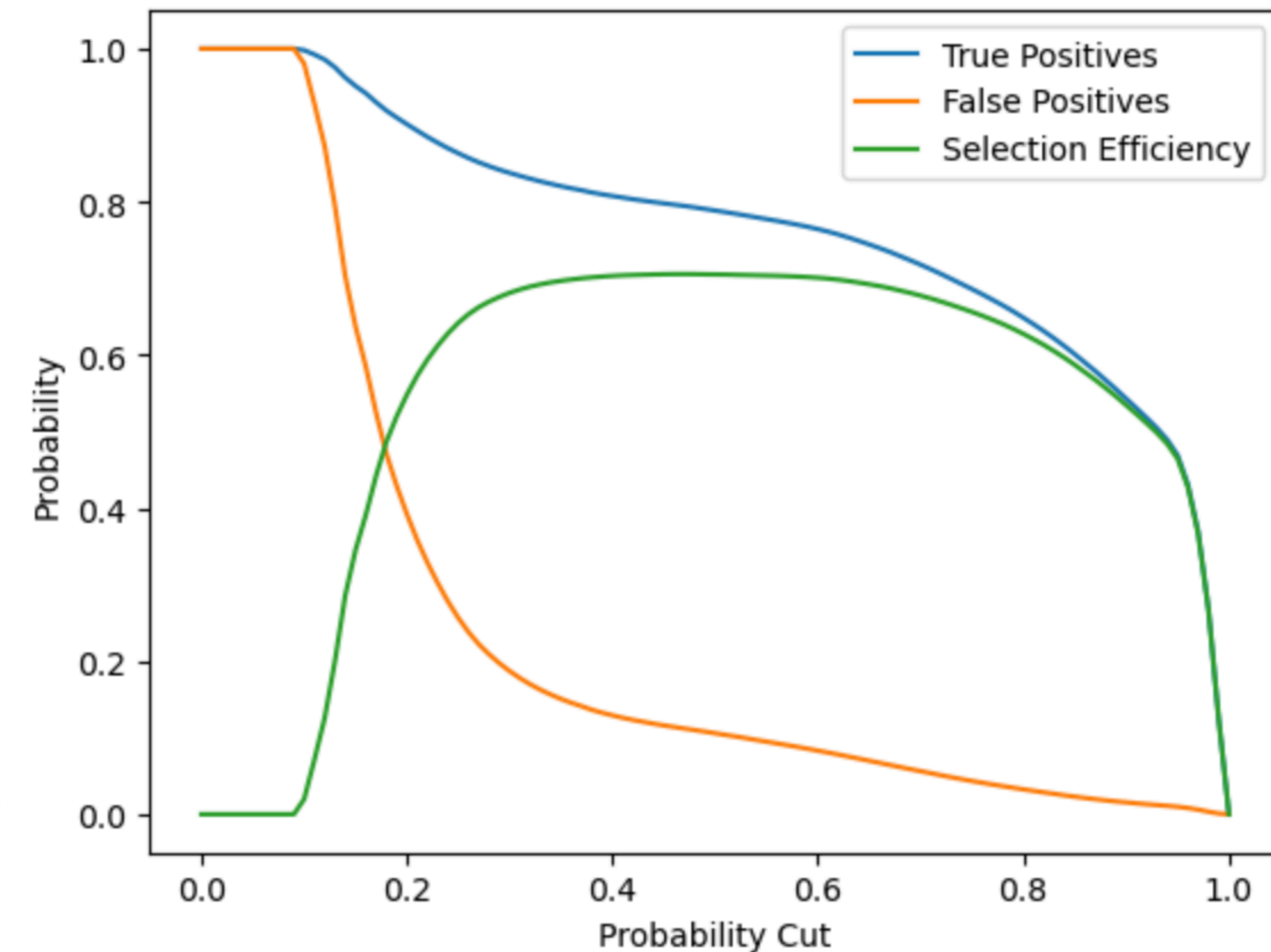
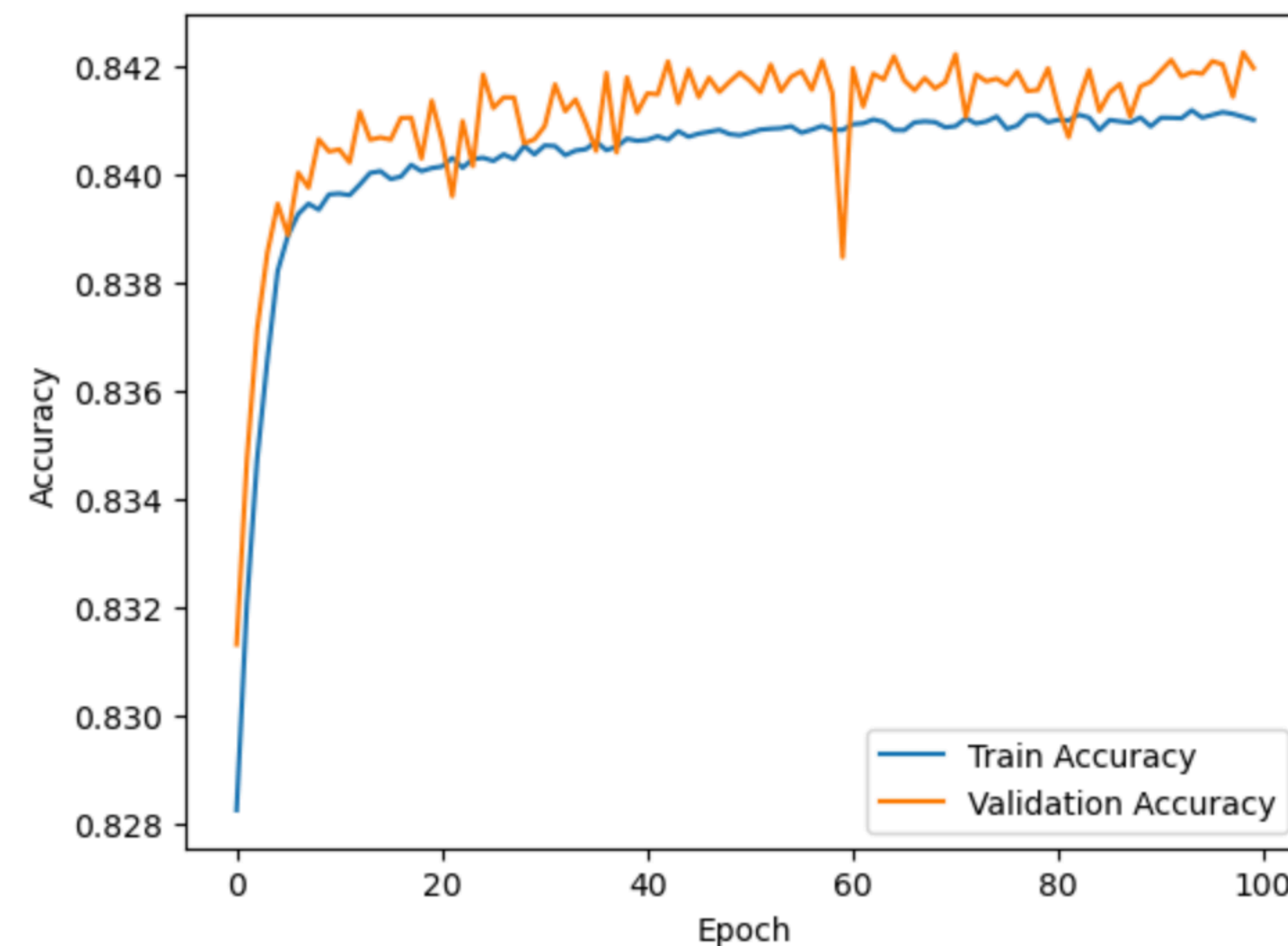
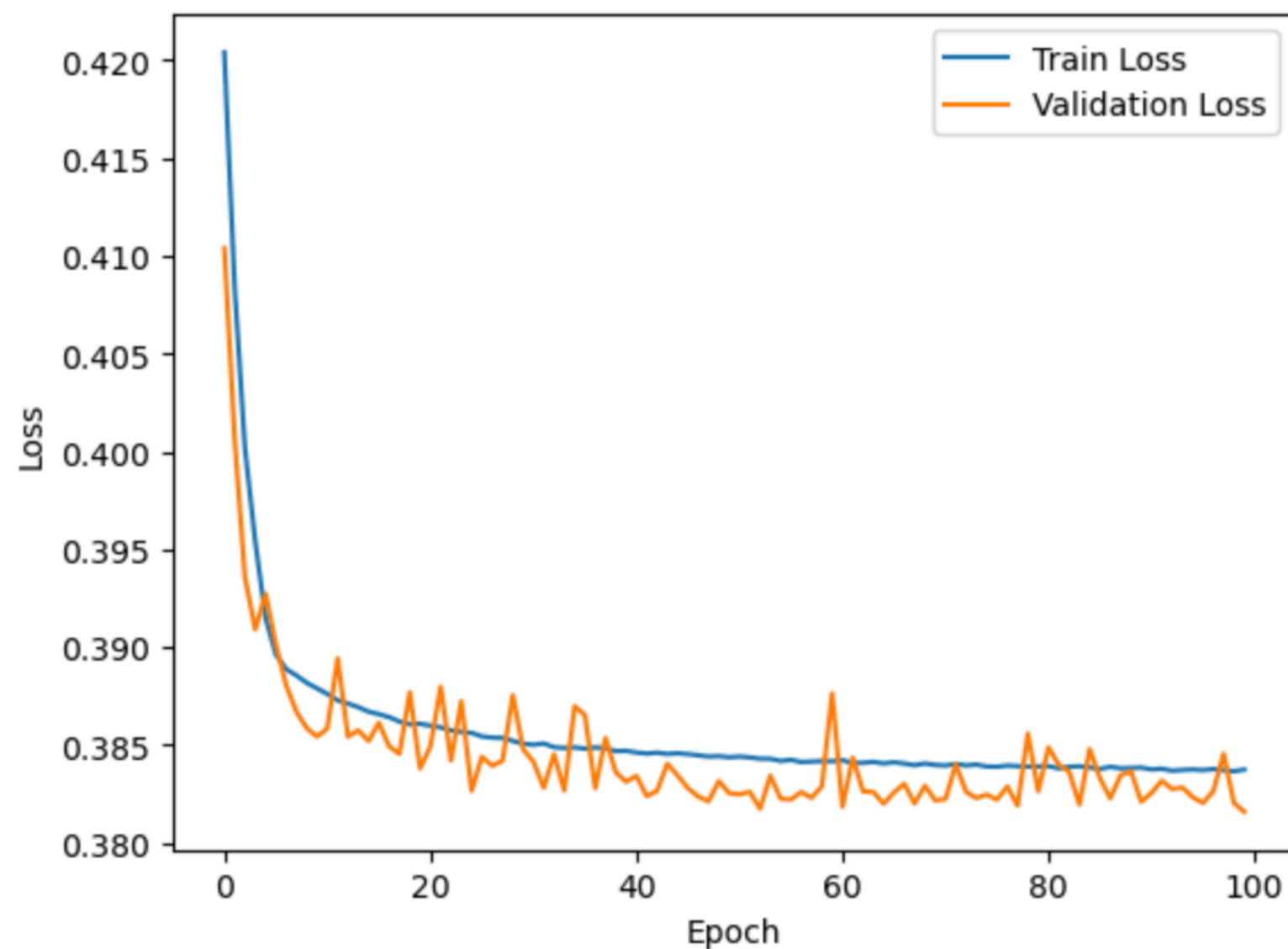


$$\text{Reaction Yield} = \text{Number of Signals} / \text{Selection Efficiency}$$



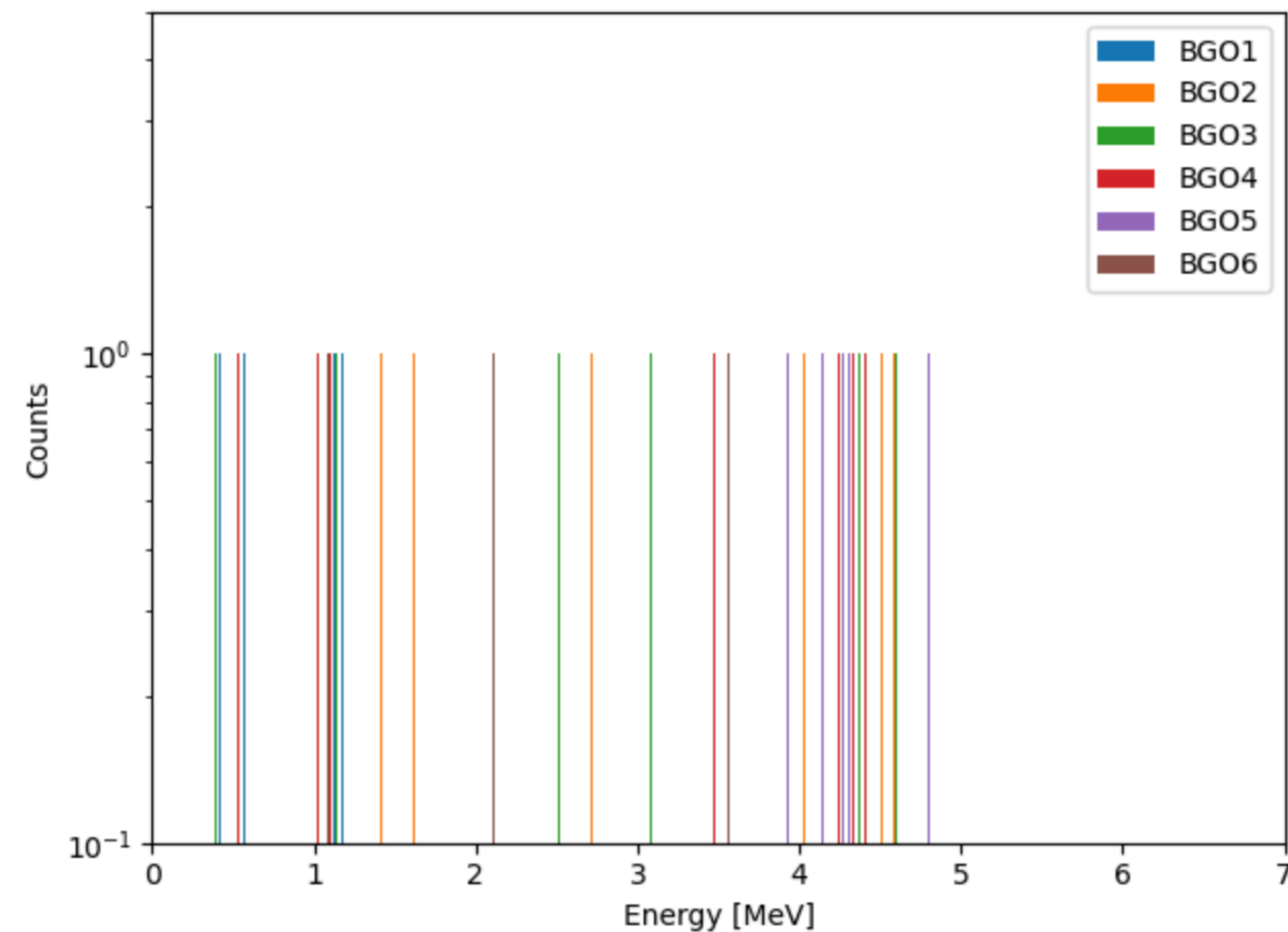
Performance (1)

- The network immediately learns the features
- The probability of **True Positives** and **False Positives** depends on the **probability cut** 
- For a cut at **0.5** we obtain an **accuracy** of **84%** and an **selection efficiency** of **70%**

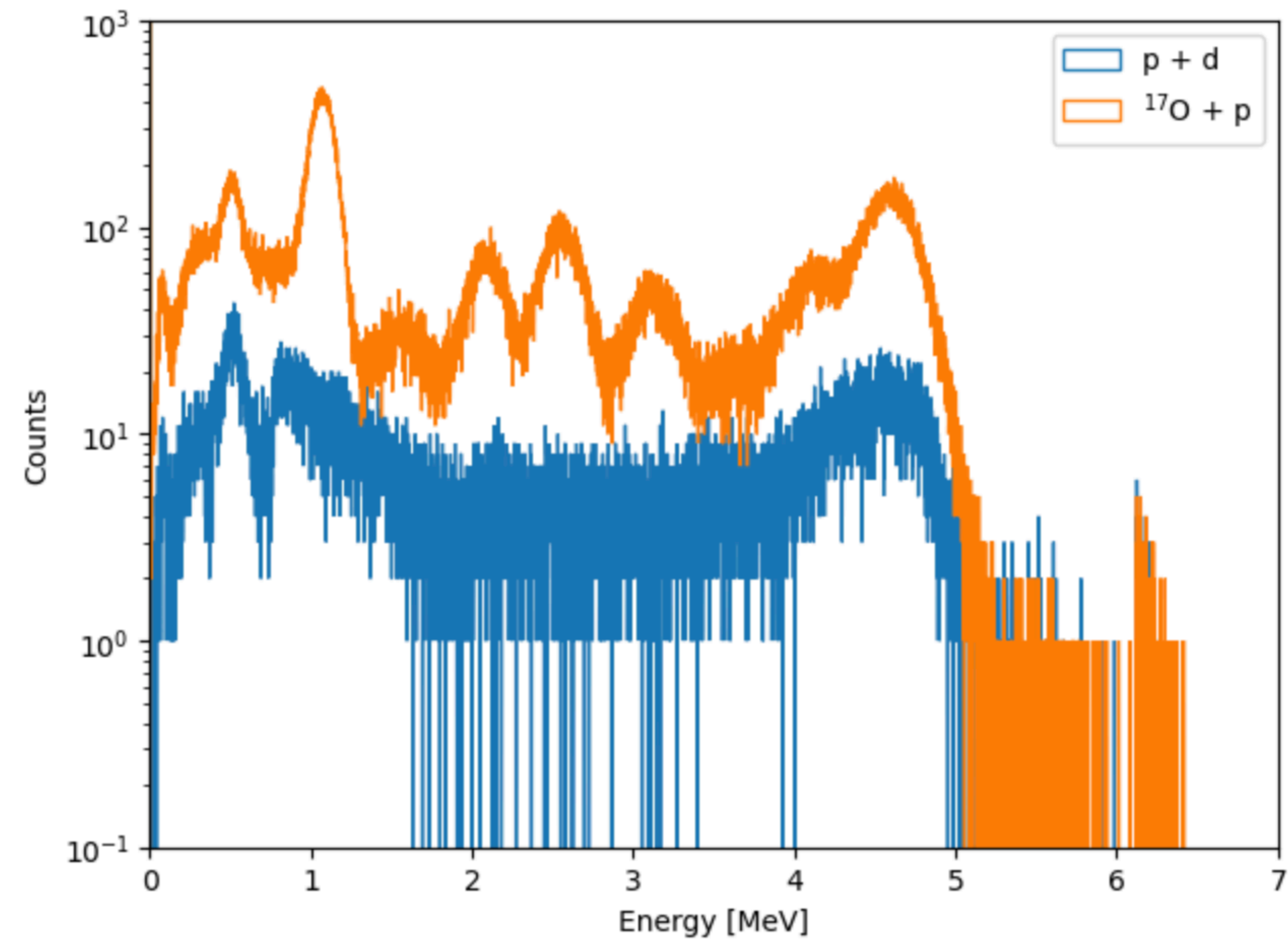


Performance (2)

True Positive Observed - All Crystals



True Positive Simulated - One Crystal



Results

Target	Counts (Classic)	Counts (NN)	Reaction Yield (Classic)	Reaction Yield (NN)
T24	41	69	102	98
T28	15	30	39	42
T44	25	50	62	71
Total	81 \pm 9	149 \pm 12	202 \pm 22	212 \pm 17

Neural Network approach
permits to **increase the
statistics and reduce the
final error!**

Classic Selection Efficiency = **40 %**

Neural Network Selection Efficiency = **70 %**



Conclusions

- Nuclear Astrophysics usually deals with **extremely low counting statistics**
- A simple **Neural Networks** can allow to obtain more **efficient cuts on the data**
- The $^{17}\text{O}(p,\gamma)^{18}\text{F}$ is the perfect example for this
- **Future Prospective:** train the NN directly on the observed γ -spectra

Thank you for attention!



LUNA Collaboration

